

DISCUSSION

Davidson: How intense are the UV carbon lines, relative to visual lines?

Stecher: In terms of H β being 1200, the combined two are 1280, approximately equivalent to the [OIII] lines.

Zuckerman: Have you an explanation for the large discrepancy between you and the Peimberts for the carbon abundance in NGC 7662? Have you or anyone else detected carbon lines in the ultraviolet from the Orion Nebula, and if so, can you derive a carbon oxygen ratio for Orion?

Stecher: The Orion Nebula has been observed by us and the strong CIII and CIV lines are not present since the temperature is much lower. We do see CII at λ 2323 and can probably get a carbon abundance from it. The Orion spectrum is primarily one of reflected starlight from the dust.

Tarter: Do your ultraviolet observations show any evidence of the MgII line (\sim 2800 Å) which is found to be very strong in quasars (which also have strong CIV and CIII lines)?

Stecher: It was just in our range and we didn't notice an image.

THE DETAILED SPECTROPHOTOMETRY OF 8 PLANETARY NEBULAE IN THE SPECTRAL REGION $\lambda\lambda$ 6,000-11,000

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The detailed spectral study in the region $\lambda\lambda$ 6,000-11,000 of eight planetary nebulae: NGC 6210, 6572, 6891, 7662, IC 2149, 4593, 4997 and BD+30° 3639 was carried out at the Crimean Station of Sternberg Astronomical Institute, using the 125-cm parabolic reflector with the contact image tube.

The absolute energy flux in emission lines was measured and used for determination of the principal physical parameters of nebulae studied: the electron temperature, T_e , the electron density, N_e , and the interstellar extinction.

The energy distribution in the merged continuum from nebula and nucleus was studied, and the separation of these two components was carried out.

The observed continuum in the studied spectral interval can entirely be explained by the Planck radiation of the nucleus and the recombination theory of planetary nebulae.