

Chemical properties of field halo and thick disk blue straggler stars: first results

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Abstract. We have observed with UVES@VLT 15 field Blue Straggler Stars for which kinematical attribution to the Halo or Thick Disk is available. We present here the first results of our detailed chemical analysis.

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1. Our program: first results

The sample of blue metal poor (BMP) nearby field stars collected by Preston *et al.* (1994) contains a large fraction of single-spectrum binaries (Preston & Sneden (2000), PS00). At least one half of them are Blue Straggler Stars (BSS) formed by the McCrea mass transfer process, that mostly works in low density environments. While PS00, working on low S/N spectra mainly aimed at deriving radial velocities, did not find any significant chemical peculiarity, subsequent dedicated studies revealed e.g., a large overabundance of Carbon and s-process elements in a few stars [for instance in CS29497-0030, see Sivarani *et al.* (2004) and Sneden *et al.* (2003)], suggesting that these stars have accreted processed material from their companions during the AGB phase.

The recent USNO-B catalog by Monet *et al.* (2003) gives proper motions for 49 stars of the PS00 sample, so we have been able to calculate the heliocentric space motions UVW, and separate Thick Disk and Halo stars on a kinematical basis (figure 1). We decided to study a fair fraction of them and derive detailed spectroscopic signatures in order i) to compare original Main Sequence and final abundances, to get information on the accreted material and the evolutionary phase(s) involved (RGB/AGB ?); ii) to find correlations with binary properties, and shed light on the formation mechanism; iii) to combine accurate abundances and abundance ratios with the space velocities and derive a better knowledge on the environment where these BMPs formed.

We observed 15 BMPs with UVES@VLT (program 72.D-0739), using a dichroic and covering the 326–451 and 665–1060 nm range at R=40000 and S/N > 100. Derivation of rotational velocities is in progress; we have already obtained abundances for the 7 stars with $V \sin i$ less than about 20 km s⁻¹. Atmospheric parameters were derived in a full spectroscopic way, using the usual minimization methods to obtain temperature, gravity and microturbulent velocity. These will have to be validated using photometric information (V and JHK); special care will be used with C-rich stars, since colors may be affected. Results of this first stage of the analysis, limited to the blue portion of our spectra, are shown in the Table.

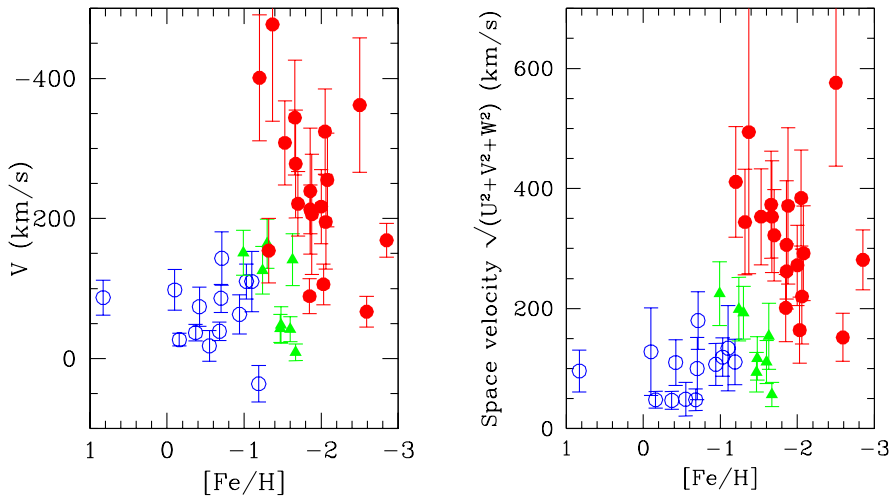


Figure 1. Rotation vector V (left) and space velocity (right) versus $[Fe/H]$ for the PS00 sample. Kinematically we can distinguish between Halo (filled dots) and Thick Disk (open dots) stars, with a few unclear cases (triangles).

Table 1. Atmospheric parameters (temperature, gravity, metallicity, microturbulent velocity), heliocentric radial velocities and iron abundances for the seven stars presently analysed

Star name	T_{eff}	$\log g$	$[A/H]$	V_t	V_r	$[Fe/H]_I$	rms	$[Fe/H]_{II}$	rms	Population
CS22966-0054	7100	3.95	-1.44	2.55	-77.5	-1.51	0.113	-1.45	0.160	Halo
CS22966-0043	7300	3.70	-2.00	3.00	-268.7	-2.04	0.180	-1.92		Halo
CS29497-0030	7150	4.28	-2.29	1.85	44.2	-2.40	0.150	-2.43	0.159	Halo
CS22956-0028	6940	4.20	-2.29	1.80	41.3	-2.31	0.163	-2.34	0.043	Halo
CS30339-0013	6900	4.05	-0.69	2.00	-30.5	-0.77	0.112	-0.76	0.101	Thick Disk
CS22894-0043	7050	4.13	-0.55	1.77	-33.5	-0.79	0.133	-0.78	0.163	Thick Disk
CS22955-0110	6650	3.75	-1.15	1.60	2.4	-1.27	0.170	-1.27	0.123	Thick Disk

Next steps in the analysis require an accurate elimination of telluric features in order to fully exploit our large wavelength coverage: several atomic lines of C, N, O, Na are present in the red, contaminated part of the spectra. N and C abundances will also be obtained from spectrum synthesis of the UV NH and CN bands and of CH band. Li appears (almost) absent from our spectra, but the visual impression has to be further checked.

Finally, we will compare our findings with existing literature, as for CS29497-0030, studied by Sivarani *et al.* (2004), Sneden *et al.* (2003) and Ivans *et al.* (this Symposium), or CS22956-0028 analysed by Lucatello (2004).

References

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