Petrogenesis of the Harsin–Sahneh serpentinized peridotites along the Zagros suture zone, western Iran: new evidence for mantle metasomatism due to oceanic slab flux

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(Received 12 March 2017; accepted 16 February 2018=first published online"34'Crtkni423:)

Abstract – The Harsin–Sahneh serpentinized peridotites are widely exposed along the Zagros suture zone in the western region of Iran and are considered to represent remnants of Neo-Tethys oceanic lithosphere at the junction of the Arabian and Iran Plates. These rocks are characterized by low contents of SiO₂ (38.8–43.5 wt%), Al₂O₃ (0.1–3.8 wt%), CaO (0.2–8.2 wt%) and TiO₂ (< 1 wt%) and high MgO contents (31.1–46.0 wt%). Their enrichments of large ion lithophile elements and light rare earth elements, with high ⁸⁷Sr/⁸⁶Sr_(i) values (0.7036–0.7109) and relatively high variations in their ϵ Nd_(t) (–7.5 to +7.8) values, indicate that the Harsin–Sahneh peridotites were metasomatized by flux released from the oceanic subducting slab in an active margin. The chemical compositions and isotopic ratios of these rocks suggest that they were formed as residue of mid-oceanic ridge basalt in the lithosphere that was then subsequently re-melted and metasomatized in a supra-subduction zone system. The occurrence of both mid-oceanic ridge and supra-subduction zone-type peridotites suggests that the heterogeneity of the upper mantle may have occurred due to the different ratios of partial melting and melt–rock reaction processes in different tectonic settings within the Neo-Tethys realm. The Harsin–Sahneh peridotites provide a good explanation of multistage melt extraction as well as melt–rock and metasomatic reactions in the mantle sequence of the Zagros ophiolite complex.

Keywords: Zagros suture zone, Neo-Tethys, metasomatic mantle, peridotite, partial melting, lithosphere

1. Introduction

Ultramafic rocks are dispersed in suture zones and orogenic belts as the basal sections of ophiolite complexes. They provide important information about melting, melt extraction processes and melt-rock interactions in the upper mantle (Zhou et al. 2005; Dilek, Furnes & Shallo, 2007; Eyuboglu, Santosh & Chung, 2011). Thus, geochemical data from peridotites and their constituent minerals can be used to characterize the origin and tectonic setting of ophiolitic rocks (Melcher et al. 2002; Karipi, Tsikouras & Hatzipanagiotou, 2006; Choi, Shervais & Mukasa, 2008; Aldanmaz et al. 2009; Pagé, Bédard & Tremblay, 2009; Pearce & Robinson, 2010; Ulrich et al. 2010). Peridotites may show surprisingly diverse variations (Parkinson et al. 1992; Parkinson & Pearce, 1998; Pearce et al. 2000), as they can range from relatively fertile peridotites to depleted restites that form after the extraction of large volumes of melt fraction (>20%) under hydrous conditions (Jean *et al.* 2010; Uysal *et al.* 2012).

The Iranian Plate is a tectonically active region within the Alpine-Himalayan orogenic belt that is squeezed between the Afro-Arabian Plate in the southwest and the Eurasian Plate in the north. It is divided into several structural zones (Fig. 1) (Stöcklin, 1968). Within the Iranian Plate, a main remnant of oceanic basement is from the Neo-Tethys Ocean. The Neo-Tethys Ocean opened from the Late Carboniferous to Permian (Berberian & King, 1981; Davoudzadeh & Schmidt, 1984; Ricou, 1994; Besse et al. 1998; Mohajjel, Fergusson & Sahandi, 2003), as Cimerian continental fragments rifted from the northern margin of Gondwana. Oceanic crust developed along this fracture, and the subduction of this oceanic crust began during the Late Triassic to Early Jurassic (Berberian & King, 1981; Davoudzadeh & Schmidt, 1984; Ricou, 1994; Besse et al. 1998; Mohajjel, Fergusson & Sahandi, 2003; Davoudian et al. 2016). The Zagros orogenic belt is related to the opening and closure of the Neo-Tethys, and is divided into three units (Fig. 1): (1) the Urmia–Dokhtar magmatic arc (Fig. 1)

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Figure 1. (Colour online) Simplified geological map of Iran (modified from Stöcklin, 1968).

was interpreted to be the product of the subduction of the Neo-Tethys oceanic slab beneath the central Iranian Plate. It records abundant Cenozoic magmatism, which is dominantly related to Eocene arc evolution (Berberian & King, 1981; Shahabpour, 2007; Omrani et al. 2008; Ghorbani & Bezanjani, 2011; Verdel et al. 2011; Ghorbani, Graham & Ghaderi, 2014). (2) The Sanandaj–Sirjan zone (SNSZ), which extends from SE to NW Iran (Fig. 1), was an active margin from the Middle Jurassic to Late Cenozoic and includes many granitoid intrusive bodies (Baharifar et al. 2004; Agard et al. 2005; Esmaeily et al. 2005; Davoudian et al. 2008; Aghazadeh et al. 2010; Azizi et al. 2011a, 2015a; Mahmoudi et al. 2011; Azizi, Zanjefili Beiranvand & Asahara, 2015b; Azizi, Kazemi & Asahara, 2017). (3) The Zagros fold-thrust belt, which consists of dismembered ophiolites and coloured mélanges, crops out along the Zagros suture zone at Neyriz and Esfandaghah in SW Iran and at Kermanshah and Piranshahr in W and NW Iran (Fig. 1). The ophiolites in the Zagros suture zone comprise relicts of Neo-Tethys oceanic lithosphere, which were finally emplaced during the closure of the ocean (Delaloye & Desmons, 1980; Desmons & Beccaluva, 1983; Ghazi & Hassanipak, 1999; Dilek & Ahmed, 2003; Alavi, 2004; Allahyari et al. 2010 Azizi et al. 2013; Saccani et al. 2013; Shafaii Moghadam & Stern, 2015; Nouri et al. 2016, 2017). The ophiolites in these belts are abundant and comprise various dismembered mafic-ultramafic assemblages that have suffered multiple phases of alteration and deformation. They display evidence of a variety of tectonic environments, ranging from a low degree of partial melting at a midoceanic ridge setting to a high degree of partial melting in an arc-related environment (Delaloye & Desmons, 1980; Ghazi & Hassanipak, 1999; Parlak, Höck



Figure 2. (Colour online) Simplified geological map of western Iran: Harsin area modified from Shahidi & Nazari (1997), Miyanrahan area from Rafia & Shahidi (2006), Kamyaran area from Sadeghian & Delvar (2006), Kermanshah area from Karimi Bavandpur & Hajihoseini (1999) and Sonqor area from Eshraghi, Jafarian & Eshraghi (1996). The surrounding broken line shows the Harsin area, and sampling locations are indicated.

& Delaloye, 2002; Dilek & Ahmed, 2003; Alavi, 2004; Allahyari *et al.* 2010, 2014; Uysal *et al.* 2012, 2016; Azizi *et al.* 2013; Saccani *et al.* 2013; Saka *et al.* 2014; Shafaii Moghadam & Stern, 2015; Ao *et al.* 2016; Nouri *et al.* 2016, 2017).

The Harsin-Sahneh ophiolite complex in eastern Kermanshah contains spectacular ophiolites and is located between the Arabian and Iranian Plates. It mainly includes ultramafic and gabbroic rocks, with minor felsic rocks (Fig. 2). Allahyari et al. (2010) demonstrated that the depleted peridotites in the Harsin-Sahneh area were generated during oceanic spreading that formed an oceanic arc basin. Some researchers (Delaloye & Desmons, 1980; Desmons & Beccaluva, 1983; Ghazi & Hassanipak, 1999; Saccani et al. 2013) have suggested that the Harsin–Sahneh ophiolite formed at a mid-ocean ridge or in island arc basins, whereas others (Wrobel Daveau et al. 2010; Azizi et al. 2013; Saccani et al. 2013; Nouri et al. 2016, 2017) have argued that this zone formed in an undeveloped oceanic basin that was affected by a mantle plume. In the current study, we present min-

eral chemistry data, whole-rock geochemical and isotopic data and geochronological data from serpentinized peridotites in the Harsin-Sahneh area along the Zagros suture zone in the eastern region of Kermanshah to identify the origin of the ultramafic rocks by comparing them with previously published data from other ultramafic rocks. The origin and tectonic affinity of the Harsin-Sahneh peridotites (HSP) are unclear because few studies (Shahidi & Nazari, 1997; Allahyari et al. 2010) have been performed on the HSP rocks and their original mineral assemblages have almost been eliminated due to the effects of serpentinization and tectonic activity. These new data are used to evaluate the formation processes (i.e. partial melting with subsequent metasomatism) and melt-rock interaction processes of the ophiolitic mantle peridotites in the Harsin–Sahneh area.

2. Regional geology and field relationships

The Harsin-Sahneh ophiolite complex is located between the cities of Harsin and Sahneh in eastern

Kermanshah, which is located in the western region of Iran (Fig. 2). This area mainly consists of Biston sedimentary rocks, Cretaceous ophiolitic rocks, Eocene gabbro and volcano-sedimentary rocks and Late Cretaceous granites (Braud, 1978; Delaloye & Desmons, 1980; Desmons & Beccaluva, 1983; Shahidi & Nazari, 1997; Ghazi & Hassanipak, 1999; Allahyari *et al.* 2010; Saccani *et al.* 2013; Whitechurch *et al.* 2013; Nouri *et al.* 2016, 2017).

The oldest rocks in this region are the Late Triassic to Cretaceous sedimentary rocks in the Biston area (Shahidi & Nazari, 1997). This body shows rough topography and is mainly massive, but it occasionally exhibits weak layering in some parts. It includes very thick crystalline limestone that crops out in most regions of the study area. The Cretaceous ophiolite complex comprises ultramafic, mafic and sedimentary mélange complexes that are divided by faults. It is difficult to clearly separate these rocks because they have been deformed and mixed with sediments. In most areas, dynamic deformation has affected the entire ophiolite complex; evidence of this deformation occurs as irregular blocks or bands within Eocene mafic and Miocene sedimentary rocks. The mafic rocks have been divided into two groups based on their intrusive age and tectonic setting: (1) Pillow basalt and coarse-grained gabbro that are black to white in colour represent Cretaceous ophiolites. Some researchers have suggested that these rocks have mid-ocean-ridge basalt (MORB) or/and within-plate origins (Allahyari et al. 2010; Allahyari, Pourmoafi & Khalatbari-Jafari, 2012). Ao et al. (2016) reported an age of 79.3 Ma for the rodingitic gabbros in this area. (2) The Eocene fine- to coarse-grained gabbros represent associated volcanic rocks. Braud (1978) suggested an Eocene age for some of the gabbroic bodies in the Harsin area. Whitechurch et al. (2013) and Ao et al. (2016), based on K-Ar and U-Pb age dating, reported ages of 56 Ma and 37 Ma for these gabbros in the northern Harsin area. These rocks are green to black in colour and exhibit rough morphology in outcrops. Our field observations show the gabbroic and sedimentary rocks were mélanged and in some areas the gabbroic and sedimentary rocks show pseudo-layer structures (Nouri, 2016). This group has been incorrectly mapped as layered gabbro on the Harsin map (Shahidi & Nazari, 1997). The trace element chemical compositions of these rocks, as well as their Sr and Nd isotopic ratios show they have affinities to an extensional tectonic regime due to the upwelling of metasomatized mantle after the Late Cretaceous collision in the Harsin area (Wrobel Daveau et al. 2010; Azizi et al. 2013; Saccani et al. 2013; Nouri et al. 2016, 2017). New field works by Wrobel Daveau et al. (2010) show the presence of detachment faults over mantle rocks that associate to upwelling of mantle. Furthermore, Braud (1978), Whitechurch et al. (2013) and Nouri et al. (2017) proposed that the body formed in an Eocene arc to back-arc setting.

The serpentinized peridotites are characterized by flat morphology; they form elevations or swelling hills that are supported by scarce vegetation (Fig. 3a). This massif chiefly includes serpentinized harzburgite and lherzolite with lenses of dunite that vary in size. Serpentinite bodies exhibit faulted and sheared contact relationships where they are exposed. Sometimes, relic lenses are schistose serpentinites and they are brecciated. Carbonate assemblages are locally developed close to the shear zone between the serpentinite and the ophiolite mélange. Farther from the shear zone, the protolith of serpentinite can be easily recognized as harzburgite, in which relicts of primary orthopyroxenes are still preserved in the serpentinized background. Serpentinized dunites are observed as tectonic slices within gabbro and limestone (Fig. 3b, c). The serpentinized peridotites are always sandwiched between Eocene and Miocene sediments and are bounded by faults (Fig. 3d, e). The cavities and fractures in the HSP rocks are filled by silica fluid that was derived from the late-stage injection of granitic magma during the Late Cretaceous (Nouri et al. 2016). The peridotites exhibit faulted relationships with the rodingitic gabbros (Fig. 3f). The rodingitic gabbros formed by a metasomatic mechanism that involved the loss of Si and alkalis from a mafic protolith during or after the serpentinization of adjacent ultramafic rocks (Nouri & Azizi, 2015) (Fig. 3g). Some peridotites exhibit faulted relationships with felsic rocks (Fig. 3h) and occur as isolated bodies and slices within the felsic bodies. They are locally cut by felsic rocks (Fig. 3i) of Late Cretaceous age (Nouri et al. 2016), thus indicating that the HSP bodies are older than the felsic dikes.

3. Petrography

The Harsin–Sahneh peridotites (HSP) are mostly serpentinized and display porphyroclastic to mesh textures. The porphyroclasite to mesh textures indicate that the rocks are tectonites. They mainly comprise olivine, orthopyroxene, scarce clinopyroxene and spinel with varying amounts of serpentinized minerals. The deformed minerals display undulose extinction and kink-band fractures; these textural features are evidences of mantle metamorphism.

Olivine generally forms elongated porphyroclasts that have been partially replaced by serpentine and show a typical mesh texture (Fig. 4a). Olivine neoblasts occur around orthopyroxene grains (Fig. 4b); the olivine neoblasts are evidence of annealing. Orthopyroxene porphyroclasts exhibit strongly corroded boundaries with embayments that have been filled by olivine and spinel. They also exhibit clinopyroxene exsolution lamellae (Fig. 4c). Orthopyroxene is frequently replaced by bastite serpentine, which is dark brown to black in colour (Fig. 4b, d). Spinel inclusions are observed in some orthopyroxene and olivine porphyroclasts (Fig. 4d). In addition, there are minor primary clinopyroxene porphyroclasts with irregular rims, which represent the residues left after partial

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Figure 3. (Colour online) Photographs of HSP bodies from the Harsin area. (a) Serpentinized peridotites with smooth morphology. (b) Serpentinized dunites as tectonic slices within limestone. (c) Serpentinized dunites as tectonic slices within gabbro. (d) Irregular blocks of serpentinites within Eocene sedimentary rocks. (e) Irregular blocks or bands of peridotites within Miocene sedimentary rocks. (f) Fault relationship between serpentinized HSP rocks and rodingitic gabbro. (g) Hydrothermal alteration in the HSP rocks. (h) Fault relationship of the HSP rocks with felsic rocks, gabbro and Miocene sediments. (i) Felsic dikes in the HSP rocks.

melting (Fig. 4e, f) (Zhou et al. 2005; Uysal et al. 2012). Secondary clinopyroxene commonly occurs as an interstitial phase along the grain boundaries of olivine (Fig. 4a, g) and orthopyroxene. This indicates that the formation of clinopyroxene occurred during the last stage, and this type of clinopyroxene is interpreted to represent the product of the crystallization of metasomatic melts (Nicolas & Prinzhofer, 1983; Luguet et al. 2001; Alard et al. 2005; Morishita et al. 2007; Uysal et al. 2015). Amphibole tends to occur as irregular growths along the rims of clinopyroxene crystals. Accessory spinels exhibit anhedral to subhedral shapes; they are reddish-brown in colour and are randomly distributed in the serpentinized peridotite (Fig. 4h). They have been affected by deformation and show elongated grain shapes with incomplete ferrian chromite rims along their grain boundaries and fractures. In some samples, the spinel-orthopyroxene association shows a symplectic texture (Fig. 4h, i).

The relatively high abundances of relict olivine (\sim 50–60%) and orthopyroxene (\sim 20–35%) compared to the low abundances of pseudomorphic clinopyroxene relicts (\sim 3–5%) probably suggest that most of the HSP rocks originated from Cpx-

harzburgites to harzburgites, as well as rare lherzolites and dunites.

4. Analytical techniques

Representative samples from the serpentinized HSP rocks were selected for chemical analyses based on their locations. The powdered samples were prepared and analysed for their major and trace element compositions. The whole-rock major element contents were determined using the wavelength dispersive Xray fluorescence (WD-XRF) technique with a Rigaku ZSX Primus II at Nagoya University. For major element analyses, 0.5 g of each rock powder sample was mixed with 5.0 g of lithium tetraborate, and the mixture was melted at 1200 °C for 12-17 min with a highfrequency bead sampler to obtain a glass bead for XRF analysis. The whole-rock Cr and Ni contents were also determined using WD-XRF, and a glass bead was prepared for trace element analysis from a mixture of 1.5 g of sample powder and 5.0 g of lithium tetraborate. The loss on ignition (LOI) of the sample was calculated based on the weight difference after ignition at 950°C.



Figure 4. (Colour online) Thin-section images of the HSP rocks from the ophiolite. (a) Peridotite showing mesh texture and remnants of olivines (XPL). (b) Peridotite showing porphyroblastic textures. Olivine neoblasts occur around orthopyroxene grains (XPL). (c) Orthopyroxene porphyroclasts showing clinopyroxene exsolution lamellae (XPL). (d) Spinel inclusions observed in some orthopyroxene porphyroclasts (PPL). (e, f) Clinopyroxene porphyroblasts with irregular rims (XPL). (g) Secondary clinopyroxene occurring as an interstitial phase along grain boundaries with olivine. (h, i) Accessory spinels exhibiting anhedral to subhedral shapes with reddishbrown colour. Spinel–orthopyroxene association with a symplectic texture is observed (PPL). OI: olivine; Srp: serpentine; Opx: orthopyroxene; SpI: spinel; Cpx: clinopyroxene (Whitney & Evans, 2010).

The HSP rocks contain low concentrations of rare earth elements (REE). Thus, 400-500 mg of each powdered sample was completely decomposed in 3 ml of HF (38%) and 0.5-1 ml of HClO₄ (70%) in a covered PTFE beaker at 120-140 °C on a hotplate in a clean room. The dissolved sample was then dried at 140 °C on a hotplate using infrared lamps. After drying, >10 ml of 2–6 M HCl was added to the dried sample to dissolve it, and the sample solution was moved to a polypropylene centrifuge tube to separate the residue from the clear upper portion. The residue was moved into a smaller PTFE vessel and then treated with $HF + HClO_4$ in a steel-jacketed bomb to ensure its complete dissolution. After the second HF decomposition, the residue fraction was mixed with the clear upper fraction, and the resulting solution was divided into two parts in a ratio of 1:20. The former was used for the quantitative analysis of trace elements, and the latter was used for Sr and Nd isotopic analysis. The former was dried and dissolved in 15-20 ml of 2% HNO₃, and its concentrations of trace elements (including REEs) were analysed using an inductively coupled plasma mass spectrometer (ICP-MS) (Agilent 7700x) at Nagoya University. The latter was dried and dissolved in 15 ml of 2.4 M HCl. To isolate and purify Sr and REE, including Nd, from the large amount of matrix elements in >400 mg of rock sample, cation exchange column separation was carried out twice. First, after loading the sample solution (15 ml), matrix elements such as Fe, Mg and Ca were eliminated by a cation exchange column (AG50W-X8, 200-400 mesh) with an eluent of 2.4 M HCl, and both Sr and REE were collected together using an eluent of 6 M HCl. The collected Sr and REE fraction was dried and dissolved in 3 ml of 2.4 M HCl. The Sr and REE solution was loaded onto the same cation exchange column with eluents of 2.4 and 6 M HCl to isolate and collect Sr and REE, respectively, in each vessel. The REE fraction was dried, then neodymium was isolated from the REE fraction using a cation exchange column with an eluent of α -hydroxyl isobutric acid (α-HIBA).

The isotopic ratios of Sr and Nd were determined using the VG Sector 54-30 and GVI IsoProbe-T thermal ionization mass spectrometers (TIMS) at Nagoya University. Mass fractionation was corrected using values of 86 Sr/ 88 Sr = 0.1194 and 146 Nd/ 144 Nd = 0.7219. The NIST-SRM987 and JNdi-1 standards (Tanaka *et al.* 2000) were used as the natural Sr and Nd isotopic ratio standards.

For some of the samples, their elemental concentrations were determined using the isotope dilution (ID) technique after adding the prepared isotope spike solutions for Rb, Sr, Sm and Nd. To isolate Rb, Sr and REE, conventional column chemistry was carried out using cation exchange resin (Bio Rad AG50W-X8, 200– 400 mesh) and an eluent of HCl. The abundances of Rb, Sr, Sm and Nd were measured using the Finnigan MAT Thermoquad THQ thermal ionization quadrupole mass spectrometer at Nagoya University.

Additionally, the trace element concentrations (including REE) of five samples (SH1, SH2, SH3, TB2, TM1) were analysed using inductively coupled plasma mass spectrometry (ICP-MS) at ALS CHEM Analytical Laboratories, North Vancouver, Canada.

After observations of the polished thin-sections were obtained, the chemical compositions of fresh minerals were determined using a JXA-8800R electron microprobe analyzer (EMPA) at Nagoya University. The acceleration voltage and beam current were set at 15 kV and 12 nA, respectively. Natural and synthetic standards were used for calibration.

5. Mineral chemistry

5.a. Spinel

Spinel grains are usually fresh and preserve their primary composition. The results of the electron microprobe (EMP) analyses of spinel in the HSP rocks are presented in Table 1a. Individually, there is no compositional variation or zoning in the spinels of the HSP rocks. These spinels are characterized by a narrow range of Cr_2O_3 contents of 35.3–42.3 wt% and Al_2O_3 contents ranging from 26.7 to 31.9 wt%. The Cr# varies from 0.44 to 0.52. The TiO₂ contents of spinels are low (<0.1 wt%), which is a common feature of ophiolitic ultramafic rocks (Arai & Yurimoto, 1994; Ahmed & Habtoor, 2015).

The minerals plot in the field of depleted peridotite (Fig. 5a) on the TiO₂ versus Cr# diagram (Dick & Bullen, 1984). The magnesium number (Mg#) values of the spinels in the HSP rocks (47.4–62.1) are generally higher than those in abyssal peridotites and SW Puerto Rico peridotites (i.e., mid-oceanic ridge rocks), and the Cr# values of the spinels are lower than those of boninites (Fig. 5b). The spinels of the HSP rocks plot close to the back-arc and eastern Cuba peridotite (i.e. back-arc basin) fields (Fig. 5b). Low Cr# (<0.6) values are typical indicators of oceanic ophiolites, including back-arc basin ophiolites, whereas arc-related ophiolitic spinels have Cr# values that are greater than 0.6 (Dick & Bullen, 1984).

5.b. Pyroxene

The results of the chemical compositions of the pyroxenes from the HSP rocks are listed in Table 1b. All analysed pyroxenes are chemically homogeneous.

The orthopyroxenes in the HSP rocks mainly comprise enstatite (En₈₀₋₉₀, Wo₁₋₁₄ and Fs₆₋₉), according to the pyroxene classification diagram (Fig. 5c) (Morimoto, 1988). Their Mg# values range from 91.3 to 95.3. These orthopyroxenes are characterized by narrow ranges of Al₂O₃ (2.5–3.6 wt %) and CaO contents (0.6–3.5 wt %), with low TiO₂ contents (<0.1 wt %). The orthopyroxenes show lower Al₂O₃ contents and higher Mg# values than MORB-type peridotites from SW Puerto Rico, and have similar compositions to back-arc peridotites from eastern Cuba (Fig. 5d).

The chemical compositions of representative clinopyroxene grains are also given in Table 1b. The residual clinopyroxenes are generally diopside in composition on the En-Wo-Fs ternary diagram (Morimoto, 1988). The Mg# values of clinopyroxene are positively correlated with their degree of depletion and degree of melting (Zhou et al. 2005; Uysal et al. 2012); clinopyroxenes in highly depleted peridotites are characterized by high Mg# values. The Mg# values of residual clinopyroxenes range from 90.7 to 92.1 (reflecting high degrees of partial melting) and are higher than those of the secondary clinopyroxenes with diopside compositions (Fig. 5c) that occur as an interstitial phase (i.e. Mg# values ranging from 84.1 to 86.6, reflecting melt-rock interactions). On the Al₂O₃ versus TiO_2 diagram (Fig. 5e), the clinopyroxenes plot in the field corresponding to a back-arc tectonic setting.

5.c. Olivine

The chemical compositions of olivines are shown in Table 1c. Both porphyroclastic and neoblastic grains show similar and homogeneous chemical compositions. The olivines in these samples plot in both the chrysolite and forsterite fields (Fig. 5f) on the olivine classification diagram (Deer, Howie & Zussman, 1992). The average Fo value in olivine reflects the fertility of a sample, and is related to the conditions of partial melting (i.e. the pressure and degree of melt extraction) (Arai, 1994; Uysal et al. 2012). The olivine porphyroclastic grains in the HSP rocks have variable Fo numbers ranging from 91.1 to 94.9 (reflecting a high degree of partial melting), whereas the neoblastic grains have variable Fo numbers ranging from 89.2 to 90.3 (due to melt-rock interactions). All of the olivines in the HSP rocks contain very low Ti and Cr concentrations, which are typically less than 0.1 wt %.

6. Whole-rock geochemistry

The whole-rock major and trace element compositions of the serpentinized HSP rocks are listed in Table 2. Their variable amounts of LOI values (1.0-6.8 wt %)

	(a) Location				Sahı	neh									Harr	ii					
	Sample				.HU	-21						UH.	6					ΠĤ	8		
	Mineral No.	Sp 1	Sp 2	3 3	Sp 4	sp 5	sp 6	Sp 7	Sp 8	gs 9	Sp 10	Sp 11	Sp 12	Sp 13	Sp 14	Sp 15	Sp 16	Sp 17	Sp 18	Sp 19	
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	SiO ₂	0.018	0.018	0.015	0.016	0.016	0.021	0.020	0.042	0.061	0.050	0.026	0.049	0.025	0.025	0.022	0.011	0.039	0.038	0.01	0
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	Al ₂ O ₃	0.019 29.04	0.00 28.89	0.040 28.36	27.41	0.000 29.23	0.010 29.23	0.010 26.73	29.11 29.11	28.55	27.40 27.40	0.004 28.27	0.000 29.28	28.71	0.020 28.28	31.86	27.43	0.044 27.32	27.05	26.68	-
	Cr_2O_3	36.35	36.88	36.81	37.36	35.34	35.44	38.32	39.40	39.60	40.91	40.99	39.81	40.16	40.57	36.85	41.69	41.54	41.59	42.31	
	FeO	20.40	20.38	20.82	21.86	21.42	21.46	22.37	16.29	16.18	16.65	16.03	16.05	15.95	16.23	15.64	16.05	16.09	15.86	16.51	
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	MgO	12.55	12.89	12.55	11.38	12.12	12.08	11.29	14.00	14.02	13.71	14.26	13.84	13.88	13.99	14.40	13.47	13.63	13.46	13.17	
	CaO	0.021	0.000	0.000	0.000	0.016	0.000	0.064	0.016	0.002	0.016	0.020	0.001	0.000	0.000	0.006	0.010	0.005	0.016	00.0	
	Na_2O	0.017	0.000	0.000	0.000	0.016	0.000	0.000	0.000	0.006	0.010	0.000	0.000	0.000	0.000	0.012	0.014	0.000	0.005	0.00	_
	K_2O	0.003	0.000	0.008	0.000	0.000	0.007	0.000	0.002	0.007	0.021	0.004	0.002	0.003	0.019	0.000	0.007	0.000	0.000	0.011	
$ \begin{array}{c c c c c c c c c c c c c c c c c c c $	(HO)	1.452	0.788	1.274	1.860	1.720	1.691	1.076	1.036	1.425	1.140	0.197	0.857	1.211	0.723	1.110	1.225	1.304	1.809	1.187	
CationSi0.0010.0000.0000.0000.0010.0010.0010.0010.0010.0010.001Ti0.0000.0010.0000.0000.0000.0010.0010.0010.0010.001Ti0.0000.0010.0010.0000.0000.0010.0010.0010.0010.001Ti0.0000.0010.0000.0000.0000.0010.0010.0010.0010.001Ti0.0000.0010.0000.0000.0000.0010.0010.0010.0010.001Ti0.0000.0010.0000.0000.0010.0010.0010.0010.0010.001Ti0.0010.0010.0010.0000.0010.0010.0010.0010.0010.001Ti0.0010.0010.0010.0000.0010.0010.0010.0010.0010.001Ti0.8730.8750.9120.8520.8540.9330.9350.9440.9560.9650.9650.9770.977Fe ³⁺ 0.8730.9250.9950.9950.9330.9320.9330.0320.0160.0010.016Mg0.5680.5790.5690.5740.5190.5540.5190.5190.6180.6530.6180.6160.613Mg52.353.051.848.150.250.147.460.56	Total	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	100.00	
Si 0.001 0.001 0.000 0.000 0.000 0.001 0.000 0.001 0.000 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.001 0.000 0.	Cation																				
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Si	0.001	0.001	0.000	0.000	0.000	0.001	0.001	0.001	0.002	0.002	0.001	0.001	0.001	0.001	0.001	0.000	0.001	0.001	00.0	_
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Ξ	0.000	0.001	0.001	0.000	0.000	0.000	0.000	0.000	0.001	0.001	0.002	0.000	0.000	0.000	0.000	0.001	0.001	0.001	0.00	
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Al	1.040	1.027	1.016	0.998	1.051	1.050	0.971	1.029	1.015	0.978	0.995	1.034	1.019	1.001	1.112	0.981	0.977	0.974	0.95	<u>~</u>
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Cr	0.873	0.879	0.885	0.912	0.852	0.854	0.933	0.935	0.944	0.979	0.968	0.944	0.956	0.963	0.863	1.001	0.996	1.005	1.02(
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Fe^{3+}	0.085	0.092	0.096	0.088	0.097	0.095	0.094	0.033	0.036	0.038	0.032	0.019	0.024	0.033	0.022	0.016	0.022	0.017	0.01	6
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Fe^{2+}	0.433	0.422	0.433	0.477	0.450	0.452	0.482	0.376	0.372	0.384	0.368	0.383	0.378	0.375	0.365	0.391	0.386	0.389	0.402	~
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Mg	0.568	0.579	0.569	0.524	0.551	0.549	0.519	0.626	0.630	0.619	0.635	0.618	0.623	0.626	0.636	0.610	0.616	0.613	0.599	6
$ \begin{array}{rrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrrr$	Total	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.000	3.00(
Cr# 0.46 0.46 0.47 0.48 0.45 0.45 0.49 0.48 0.48 0.48 0.50 0.49 0.48 0.49 0.48 0.49 0.44 0.50 0.50 0.51	Mg#	52.3	53.0	51.8	48.1	50.2	50.1	47.4	60.5	60.7	59.5	61.3	60.6	60.8	60.6	62.1	59.9	60.1	60.2	58.7	
	Cr#	0.46	0.46	0.47	0.48	0.45	0.45	0.49	0.48	0.48	0.50	0.49	0.48	0.48	0.49	0.44	0.50	0.50	0.51	0.52	

Table 1. Chemical compositions of (a) spinel, (b) pyroxene and (c) olivine, based on EMP analyses

Table 1. Conti	nued																	
(q)																		
Location					Sahı	neh								Hars	in			
Sample			-HU	-21				-HU	×					-HU	6-			
Mineral No.	Opx 1	Opx 2	Opx 3	Opx 4	Opx 5	0px 6	Opx	Opx 8	0px 9	Opx 10	Opx 11	Opx 12	Opx 13	Opx 14	Cpx 15	Cpx 16	Cpx 17	Cpx 18
SiO ₂ TiO ₅	56.41 0.047	56.98 0.023	56.97 0.030	56.45 0.042	56.30 0.000	56.00 0.014	56.04 0.000	56.87 0.036	56.24 0.010	56.01 0.016	57.03 0.000	55.77 0.001	56.84 0.046	57.06 0.011	52.00 0.385	52.74 0.258	50.85 0.308	52.05 0.313
Al_2O_3	3.574	2.928	3.021	3.063	3.045	3.006	3.035	3.073	2.912	2.837	2.942	2.504	2.962	2.997	3.881	3.974	3.899	3.993
Cr_2O_3	0.996	0.924	0.851	0.901	1.015	0.848	0.910	0.936	0.880	0.915	0.763	1.004	0.874	0.956	1.198	1.115	1.223	1.258
FeU MaO	4.015 31 82	23.77	4.235	4.945 34.00	4.627	4.94 33.21	4.945 33.76	4.669 37 08	4.989 34.01	5.009 34 36	24.18 34.18	5.308 31.62	3.954 34.43	3.394 37.08	2.602	2.426 15.82	5.072 16.74	2.805 16 31
CaO	28.10	1.923	1.939	1.151	1.745	2.515	2.315	2.840	1.077	1.042	0.649	3.504	0.863	2.826	23.96	23.89	22.73	22.91
Na ₂ O	0.629	0.046	0.221	0.019	0.020	0.042	0.027	0.031	0.015	0.000	0.000	0.048	0.000	0.029	0.457	0.414	0.324	0.484
K20	0.009	0.000	0.000	0.013	0.000	0.008	0.008	0.000	0.000	0.000	0.000	0.001	0.015	0.000	0.014	0.005	0.013	0.015
Iotal	100.30	72.66	100.45	100.59	100.30	100.59	100.54	101.43	100.13	100.18	100.76	c/.66	99.98	62.00	18.66	100.64	99.16	100.14
Cation Si	1.937	1.963	1.952	1.932	1.933	1.920	1.923	1.937	1.932	1.921	1.948	1.941	1.950	1.960	1.900	1.909	1.864	1.890
Ti	0.001	0.001	0.001	0.001	0.000	0.000	0.000	0.001	0.000	0.000	0.000	0.000	0.001	0.000	0.011	0.007	0.008	0.009
Al	0.145	0.119	0.122	0.124	0.123	0.121	0.123	0.123	0.118	0.115	0.118	0.103	0.120	0.121	0.167	0.170	0.168	0.171
Cr	0.027	0.025	0.023	0.024	0.028	0.023	0.025	0.025	0.024	0.025	0.021	0.028	0.024	0.026	0.035	0.032	0.035	0.036
Fe^{3+}	0.000	0.000	0.000	0.000	0.000	0.017	0.009	0.000	0.000	0.017	0.000	0.000	0.000	0.000	0.010	0.000	0.075	0.029
Fe ^{≠+} Ma	011.0	0.080	0.121	0.142	0.133	0.125	0.133	0.133	0.143	0.120	0.148	0.154 1.640	0.113	0.098	0.0/0	0.0/3	0.019	000.0
Ca	0.103	0.071	0.071	0.042	0.064	0.092	0.085	0.104	0.040	0.038	0.024	0.131	0.032	0.104	0.938	0.927	0.893	0.892
Na	0.042	0.003	0.015	0.001	0.001	0.003	0.002	0.002	0.001	0.000	0.000	0.003	0.000	0.002	0.032	0.029	0.023	0.034
Total	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000	4.000
Mg#	93.4	95.3	93.3	92.5	93.0	92.3	92.3	92.6	92.4	92.4	92.4	91.3	93.9	94.5	91.3	92.1	90.7	91.2
End-member																		
Wo	5.58	3.76	3.77	2.20	3.35	4.78	4.41	5.42 87.6	2.06 00 5	1.98	1.24	6.79 of 7	1.66	5.50	50.56	49.99	46.95 48 1	47.93
En Fe	00.2 6 74	91.7 4 54	09.0 6.43	7 38	6 94	01.9 733	2.00 7.36	07.0 6.06	20.0 7 45	7.41	0.14 7 76	8 07	72.4 5.05	69.5 5 16	4.70 4.70	40.1 3 06	40.1 4 06	4 58
Wo=Wollastor	ite; En=Er	nstatite; Fs-	=Ferrosilite		1.0	00.1	00.1	00		11.1	0	70.0	0.0	01.0	(7·L	0		07.H

Location				Sahr	leh									Har	sin					
Sample						HU	-21									HN	6			
Mineral No.	Cpx 1	Cpx 2	Cpx 3	Cpx 4	Cpx 5	Cpx 6	Cpx 7	Cpx 8	Cpx 9	Cpx 10	Cpx 11	Cpx 12	Cpx 13	Cpx 14	Cpx 15	Cpx 16	Cpx 17	Cpx 18	Cpx 19	Cpx 20
SiO ₂ TiO ₂ Al ₂ O ₃ FeO FeO MmO MmO MmO MmO K ₂ O MmO K ₂ O Cation Total Ti Ti Ti Ti Total Mg K ₂ O Cation Total Mg Fe ²⁺ Fe ²⁺ Cation Si Total Cation End Catoo Si FeO MmO Catoo Si Si Si Si Si Si Si Si Si Si Si Si Si	$\begin{array}{c} 52.99\\ 0.435\\ 0.828\\ 0.828\\ 0.828\\ 15.77\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.118\\ 0.012\\ 0.118\\ 0.012\\ 0.118\\ 0.012\\ 0.118\\ 0.012\\ 0.012\\ 0.012\\ 0.035\\ 0.005\\ 0.$	52.50 0.448 3.200 0.724 15.80 0.137 0.137 0.137 0.137 0.013 0.0012 0.0012 0.0128 0.0138 0.021 0.002 0.146 0.864 0.864 0.864 0.864 0.864 0.864 0.864 0.864 0.864 0.864 0.864 0.031 3.139 0.031 0.001 0.031 0.0012 0.0310000000000	52.08 0.514 3.173 0.877 4.736 15.39 15.39 0.016 0.014 0.014 0.014 0.0138 0.026 0.014 0.0138 0.026 0.014 0.1146 0.0138 0.026 0.0380 0.03800 0.0380 0.0380 0.03800 0.0380000000000	52.42 0.518 3.198 3.198 0.756 4.699 15.42 0.145 0.528 n.d 1.920 0.138 0.014 0.138 0.022 0.014 0.138 0.022 0.003 0.138 0.037 0.138 0.037 1.920 0.037 1.920 0.037 1.920 0.037 1.920 1.920 0.037 1.9200 1.92000 1.9200 1.9200 1.9200 1.9200 1.92000 1.92000 1.92000 1.92000 1.92000 1.92000000000000000000000000000000000000	52.81 0.650 3.092 0.559 0.559 15.44 0.127 0.127 0.127 0.127 0.127 0.127 0.127 0.005 0.005 0.0018 0.0016 0.0016 0.0036 0.141 0.885 0.0036 0.0131 0.0141 0.0036 0.0036 0.0036 0.0036 0.0036 0.0036 0.0036 0.0036 0.141 0.0036	51.34 0.663 3.813 0.56 5.182 16.37 0.52 0.103 0.52 0.103 0.52 0.52 0.52 0.52 0.52 0.52 0.52 0.52	52.30 0.993 3.032 0.098 5.289 15.88 0.198 0.44 0.44 0.44 0.44 0.131 0.003 0.016 0.131 0.003 0.016 0.138 0.016 0.131 0.003 0.138 0.016 0.138 0.015 0.138 0.015 0.138 0.015 0.138 0.015 0.144 0.131 0.003 0.138 0.003 0.144 0.131 0.003 0.138 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.144 0.131 0.003 0.0138 0.144 0.003 0.131 0.003 0.144 0.003 0.131 0.003 0.144 0.003 0.131 0.003 0.144 0.131 0.003 0.135 0.003 0.144 0.131 0.003 0.003 0.144 0.003 0.003 0.138 0.003 0.144 0.144 0.003 0.138 0.003 0.144 0.003 0.003 0.138 0.003 0.144 0.003 0.131 0.003 0.003 0.003 0.003 0.138 0.003 0.003 0.144 0.003 0.0	52.64 0.782 3.032 5.248 15.70 115.70 0.1133 2.222 0.1133 0.004 0.001 0.001 0.004 0.001 0.001 0.001 0.002 0.001 0.002 0.0	52.30 1.013 3.114 5.211 15.49 0.151 15.49 0.151 1.5.49 0.151 1.5.49 0.151 0.151 0.151 0.028 0.028 0.028 0.028 0.028 0.037 0.037 3.105 84.1 47.0 44.6	51.51 1.149 3.553 0.119 5.139 5.139 5.139 15.41 0.0126 0.2226 0.2226 0.246 0.032 0.032 0.033 0.017 0.033 0.017 0.033 0.017 0.033 0.017 0.034 0.003 0.034 0.003 0.034 0.003 0.154 0.032 0.154 0.032 0.154 0.032 0.154 0.032 0.154 0.032 0.154 0.032 0.154 0.126 0.032 0.0177 0.034 0.00177 0.034 0.034 0.0034 0.034 0.0034 0.0034 0.034 0.0034 0.034 0.0034 0.0340000000000	52.10 1.103 3.191 0.182 5.169 15.41 0.156 0.156 0.156 0.156 0.156 0.156 0.156 0.156 0.156 0.156 0.006 1.904 0.005 0.005 0.030 0.137 0.033 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 0.125 1.125 1.125 1.125 1.125 1.125 1.125 1.125 1.125 1.125 1.137 1.137	52.47 0.775 0.148 5.014 15.70 0.161 22.53 0.005 0.002 0.003 0.004 0.003 0.004 0.003 0.004 0.004 0.003 0.004 0.003 0.002 0.003 0.023 0.023 0.028 0.0141 0.020 0.021 0.028 0.021 0.028 0.021 0.028 0.035 0.0161 0.020 0.035 0.0161 0.020 0.035 0.0161 0.020 0.035 0.020 0.000000 0.0200000000	52.86 0.933 2.663 15.43 0.308 15.43 0.187 0.187 0.187 0.187 0.187 0.187 0.026 0.115 0.026 0.115 0.009 0.115 0.009 0.115 0.000 0.1149 0.000 0.1149 0.000 0.1149 0.001 0.1149 0.001 0.1149 0.001 0.1149 0.001 0.1187 1.432 0.001 0.0187 0.038 0.187 0.038 0.038 0.036 0.036 0.036 0.187 0.036 0.031 0.036 0.036 0.031 0.036 0.031 0.036 0.031 0.036 0.031 0.036 0.0310 0.0310 0.0310 0.0310 0.0310 0.0310 0.0310 0.0310 0.0310 0.03100 0.03100 0.0310000000000	51.36 0.645 4.101 0.954 4.304 15.10 0.134 22.61 0.134 99.72 1.891 1.891 1.891 0.018 0.018 0.018 0.0178 0.0178 0.0178 0.013 0.013 0.013 0.013 104 148.1	52.20 0.630 3.918 0.844 4.294 15.47 0.137 0.137 0.137 0.137 0.137 0.137 0.137 0.137 0.137 0.016 0.016 0.017 0.016 0.005 0.123 0.840 0.879 0.024 0.005 0.123 0.840 0.017 0.016 0.017 0.017 0.017 0.016 0.017 0.017 0.017 0.016 0.017 0.016 0.017 0.016 0.017 0.016 0.017 0.016 0.017 0.016 0.016 0.017 0.016 0.017 0.016 0.017 0.017 0.016 0.017 0.	53.15 0.383 0.774 0.774 0.774 16.79 0.135 0.035 0.005 0.135 0.005 0.135 0.005 0.135 0.002 0.135 0.002 0.135 0.002 0.135 0.022 0.000 0.135 0.022 00000000	53.09 0.394 3.143 0.764 4.526 15.99 0.106 0.013 0.013 0.013 0.022 0.034 0.022 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.134 0.034 0.134 0.034 0.134 0.034 0.134 0.034 0.134 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.034 0.035 0.034 0.034 0.034 0.035	53.46 0.409 2.767 0.392 16.04 16.04 16.04 0.473 0.473 0.473 0.473 0.473 0.473 0.473 0.473 0.118 0.011 0.011 0.011 0.011 0.011 0.033 3.111 86.4 45.9	53.66 0.458 0.495 1.495 0.104 15.95 0.104 0.104 0.104 0.104 0.104 0.104 0.104 0.104 0.104 0.104 0.005 1.949 0.0013 0.014 0.013 0.014 0.013 0.013 0.013 0.013 0.014 0.013 0.014 0.015 0.014 0.014 0.014 0.014 0.014 0.014 0.022 3.090 86.7 45.4	52.76 0.681 2.9000 5.006 5.006 5.006 5.006 0.155 0.165 0.165 0.165 0.152 0.013 0.0125 0.015 0.0125 0.0125 0.0125 0.0125 0.0125 0.0125 0.0125 0.0125 0.0125 0.0125 0.0012 0.0125 0.00000000000000000000000000000000000
Fs Wo=Wollasto n.d.= not dete	7.53 nite; En=] cted.	7.83 Enstatite;	7.83 Fs=Ferro	7.70 silite.	7.57	8.46	8.54	8.51	8.41	8.41	8.34	8.11	7.93	7.15	7.07	7.45	7.35	7.20	6.95	8.14

Table 1. Continued

(c)				-																
Location				Sahneh									-	Harsın						
Sample				UH-21							UH-	8						0-HU		
Mineral Points	0 1	01 2	01 3	01 4	01 5	01 6	0I 7	01 8	01 9	01 10	01 11	01 12	01 13	01 14	01 15	01 16	01 17	01 18	01 19	01 20
$ \begin{array}{c} SiO_{2} \\ TiO_{2} \\ TiO_{2} \\ FeO_{3} \\ FeO_{3} \\ FeO_{3} \\ FeO_{3} \\ Cation \\ Total \\ Total \\ Total \\ Total \\ Fe^{3+}_{2+} \\ Fe^{3+}$	$\begin{array}{c} 41.51\\ 0.000\\ 0.000\\ 0.048\\ 10.13\\ 49.68\\ 0.048\\ 0.017\\ 0.159\\ 0.000\\ 0.$	$\begin{array}{c} 42.37\\ 0.000\\ 0.021\\ 0.001\\ 10.38\\ 49.65\\ 0.007\\ 0.007\\ 0.007\\ 0.007\\ 0.000\\ 0.$	$\begin{array}{c} 41.87\\ 0.011\\ 0.006\\ 0.036\\ 10.48\\ 49.06\\ 0.055\\ 0.055\\ 0.000\\ 0.015\\ 0.139\\ 0.100\\ 0.010\\ 0.000\\ 0.$	$\begin{array}{c} 41.60\\ 0.011\\ 0.004\\ 0.034\\ 10.12\\ 50.15\\ 0.055\\ 0.055\\ 0.055\\ 0.012\\ 0.124\\ 0.124\\ 0.124\\ 0.124\\ 0.000\\ 0.000\\ 0.000\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.000\\ 0.$	$\begin{array}{c} 41.26\\ 0.032\\ 0.026\\ 0.000\\ 8.373\\ 50.07\\ 0.074\\ 0.074\\ 0.008\\ 0.012\\ 0.001\\ 0.001\\ 0.001\\ 0.001\\ 0.000\\ 0.000\\ 0.171\\ 1.820\\ 0.000\\ 0.000\\ 0.000\\ 0.171\\ 1.820\\ 0.000\\ 0.$	$\begin{array}{c} 41.99\\ 0.027\\ 0.000\\ 0.012\\ 9.180\\ 0.039\\ 0.039\\ 0.039\\ 0.039\\ 0.039\\ 0.000\\ 0.$	$\begin{array}{c} 43.60\\ 0.0006\\ 0.1006\\ 0.0065\\ 50.67\\ 50.67\\ 0.003\\ 0.013\\ 0.013\\ 0.013\\ 0.013\\ 0.003\\ 0.000\\$	42.43 0.000 0.050 0.000 8.320 8.320 48.77 0.131 0.004 0.004 0.004 0.000 0.000 0.001 0.001 0.000 0.000 0.000 0.000 0.000 0.000 0.000 0.584 0.584 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.000000	$\begin{array}{c} 40.87\\ 0.013\\ 0.000\\ 8.472\\ 50.03\\ 0.133\\ 0.03\\ 0.03\\ 0.025\\ 0.448\\ 0.00\\ 0.00\\ 0.003\\ 0.000\\$	41.16 0.000 0.027 6.781 6.781 6.781 0.027 0.027 0.027 0.027 0.027 0.027 0.027 0.023 0.0388 0.013 0.000 0.000 0.000 0.000 0.0139 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.00000 0.00000 0.00000 0.000000	43.91 0.000 0.073 0.073 0.073 0.073 0.073 0.073 0.007 0.008 0.007 0.065 0.007 0.065 0.000 0.000 0.000 0.000 0.000 0.000 0.007 0.000 0.007 0.000 0.000 0.007 0.0000 0.00000 0.0000 0.0000 0.00000 0.00000 0.0000 0.0000 0.000	$\begin{array}{c} 40.57\\ 0.015\\ 0.029\\ 0.020\\ 0.029\\ 51.26\\ 0.0137\\ 0.016\\ 0.003\\ 0.567\\ 0.003\\ 0.000\\ 0$	40.95 0.000 0.198 0.000 8.243 49.68 0.063 0.000 0.499 0.000 0.499 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.000000	40.42 0.000 0.115 0.000 0.259 0.259 0.022 0.022 0.022 0.022 0.030 0.000 0.003 0.003 0.003 0.003 0.000 0.000 0.015 0.000 0.025 0.000 0.000 0.000 0.025 0.0000 0.00000 0.0000 0.0000 0.0000 0.00000 0.00000 0.0000 0.0000	$\begin{array}{c} 42.36\\ 0.028\\ 0.040\\ 0.012\\ 7.963\\ 0.011\\ 0.011\\ 0.011\\ 0.001\\ 0.001\\ 0.000\\ 0.$	40.26 0.000 0.087 0.000 5.833 5.833 5.833 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.0000 0.000000	$\begin{array}{c} 42.21\\ 0.012\\ 0.011\\ 0.014\\ 0.013\\ 0.014\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.031\\ 0.000\\ 0.$	40.80 0.000 0.195 0.080 6.899 6.899 6.899 0.064 0.064 0.064 0.064 0.018 0.456 0.018 0.456 0.000 0.018 0.456 0.000 0.018 0.456 0.000 0.018 0.018 0.054 0.000 0.018 0.054 0.000 0.018 0.054 0.000 0.018 0.054 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.000 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.000 0.018 0.0000 0.00000 0.00000 0.0000 0.00000 0.00000 0.00000 0.00000000	$\begin{array}{c} 40.51\\ 0.00\\ 0.000\\ 0.048\\ 0.05\\ 0.02\\ 0.017\\ 0.159\\ 0.00\\ 0.000\\$	$\begin{array}{c} 40.37\\ 0.000\\ 0.021\\ 0.009\\ 0.091\\ 0.007\\ 0.007\\ 0.000\\ 0.$
Ca Total End-member Fo Fa Fo=Forsterite:	0.001 3.000 89.7 10.3 Fa= Faval	0.002 3.000 89.4 10.5 ite	0.001 3.000 89.2 10.7	0.001 3.000 89.8 10.2	0.002 3.000 91.3 8.6	0.001 3.000 90.3 9.6	0.002 3.000 94.9 4.99	0.003 3.000 91.1 8.7	0.003 3.000 8.7 8.7	0.002 3.000 92.9 7.01	0.008 3.000 92.5 7.13	0.004 3.000 92.3 7.5	0.002 3.000 91.4 8.5	0.007 3.000 91.6 8.1	0.003 3.000 91.5 8.32	0.001 3.000 94.2 5.8	0.004 3.000 90.3 9.5	0.002 3.000 92.8 7.11	0.001 3.000 89.7 10.3	0.002 3.000 89.4 10.5

Table 1. Continued

https://doi.org/10.1017/S0016756818000201 Published online by Cambridge University Press



Figure 5. (Colour online) (a) Diagram of Cr_2O_3 versus TiO₂ in spinel. Fields for ultramafic rocks are after Herbert (1982), Jan & Windley (1990) and Zhou & Kerrich (1992). (b) Mg# versus Cr# diagram for spinel. Fields of spinel compositions in abyssal peridotite, MORB-BAB, island arc tholeiite and boninite are from Dick & Bullen (1984) and the GEOROC database. Data of spinel in eastern Cuba dunites and harzburgites are from Marchesi *et al.* (2006), and those in SW Puerto Rico Iherzolites are from Marchesi *et al.* (2011). (c) Classification of pyroxenes (Morimoto, 1988). (d) Diagram of Al₂O₃ (wt%) versus Mg# in orthopyroxene. Data of orthopyroxene in eastern Cuba dunites and harzburgite are from Marchesi *et al.* (2006), and those in SW Puerto Rico Iherzolites are from Marchesi *et al.* (2011). (e) TiO₂ + Cr₂O₃ (wt%) versus Al₂O₃ (wt%) diagram (after Hout *et al.* 2002) for clinopyroxene. Fields are outlined for clinopyroxene compositions in boninite (Van der Laan *et al.* 1992), island arc tholeiite and back-arc basalt (Hawkins & Allen, 1994) and NMORB (Stakes & Franklin, 1994). (f) Classification of olivine minerals (Deer, Howie & Zussman, 1992).

mainly indicate that all of the HSP samples have experienced moderate to low degrees of serpentinization (Table 2). They are characterized by low contents of SiO₂ (38.8-43.5 wt%), Al₂O₃ (0.0-3.8 wt%) and CaO (0.2-8.2 wt%). Al₂O₃ mainly resides in spinel and is partly incorporated in primary pyroxene (Li *et al.* 2015). The samples display variable degrees of depletion, with MgO contents ranging from 31.1 to 45.9 wt%. In addition, their trace element abundances are highly variable. These samples have variable contents of Cr (96–7560 ppm) and Ni (112–2880 ppm).

Most of their Cr and Ni contents are higher than those of the primitive mantle (Ni = 2090 ppm and Cr = 3240 ppm; Hart & Zindler, 1986), which probably reflects their high olivine/pyroxene ratios and relatively high Cr-spinel contents (Whattam, Cho & Smith, 2011).

On the MgO/SiO₂ versus Al_2O_3/SiO_2 diagram (Fig. 6), most of the peridotites plot at low Al_2O_3/SiO_2 values (<0.04). The most depleted (refractory) compositions are defined by Al_2O_3/SiO_2 and MgO/SiO₂ ratios of <0.01 and >1, respectively (Jagoutz *et al.* 1979;

Table 2. Chemical compositions of the HSP

Sample Rock type	UH-3 Lhz	UH-5 Cpx-hz	UH-8 Du	UH-9 Hz	UH-10 Hz	UH-17 Du	UH-19 Cpx-hz	UH-20 Lhz
Location	Harsin	Harsin	Harsin	Harsin	Sahneh	Harsin	Sahneh	Sahneh
SiO ₂ (wt%)	43.54	43.45	42.11	38.81	41.13	40.80	42.94	41.86
TiO ₂	0.04	0.04	0.02	0.02	0.02	0.01	0.01	0.06
Al_2O_3	1.32	1.03	0.93	3.47	3.18	0.34	1.39	3.81
Fe ₂ O ₃	11.17	9.74	9.04	9.70	7.74	9.43	9.81	12.39
MnO	0.17	0.16	0.11	0.11	0.11	0.13	0.15	0.18
MgO	33.97	37.96	44.02	44.37	38.69	45.97	42.35	31.09
CaO	4.28	5.91	0.58	0.52	1.54	0.62	1.90	5.22
Na ₂ O	0.13	0.14	0.07	0.07	0.08	0.03	0.05	0.10
K ₂ O	0.00	0.00	0.03	0.09	0.03	0.00	0.00	0.00
P_2O_5	0.007	0.007	0.014	0.017	0.014	0.006	0.007	0.007
Cr_2O_3	0.62	0.57	0.44	0.59	0.95	0.58	0.59	0.47
NiO	0.18	0.17	0.36	0.34	0.26	0.34	0.33	0.15
LOI	4.45	1.35	2.66	1.78	6.56	2.06	0.97	4.80
Total	99.87	100.50	100.32	99.75	100.06	100.31	100.48	100.13
Sc (ppm)	18.2	20.1	5.57	5.71	4.46	9.09	13.0	20.2
V	75.0	83.4	18.7	35.8	18.8	39.9	61.2	59.2
Cr	4240	3900	2980	4030	6470	4000	4010	3230
Co	95.9	87.4	99.6	108	96.5	122	106	113
Ni	1410	1340	2840	2710	2060	2670	2590	1160
Cu	15.3	15.3	2.73	319	14.0	9.19	30.0	13.7
Zn	47.3	44.1	38.6	36.9	34.4	44.0	48.0	57.5
Ga	1.82	1.73	0.614	4.44	2.76	0.67	1.64	2.81
Rb	0.559	0.670	0.912	1.82	0.425	0.460	1.24	1.17
Sr	4.59	6.49	5.67	10.8	4.18	4.96	2.20	7.90
Zr	0.834	1.30	0.845	1.95	1.31	0.195	0.164	0.436
Nb	0.076	0.105	0.082	0.158	0.057	0.019	0.038	n.d
Cs	0.022	0.022	0.036	0.119	0.013	0.006	0.018	0.029
Ba	3.36	4.22	2.00	6.38	1.33	0.497	0.632	0.656
Pb	0.147	0.241	0.164	0.372	0.120	0.066	0.128	0.106
Th	0.022	0.036	0.027	0.116	0.020	0.003	0.090	0.052
U	0.032	0.020	0.005	0.073	0.005	0.002	0.008	0.007
Ht	0.029	0.042	0.022	0.062	0.035	0.004	0.005	0.024
Ta	0.013	0.018	0.009	0.034	0.010	0.004	0.026	0.009
Y T	0.913	1.15	0.411	0.321	0.343	0.200	0.541	0.875
La (ppm)	0.092	0.161	0.320	0.41/	0.099	0.145	0.529	0.289
Ce D	0.166	0.298	0.500	0.880	0.236	0.200	0.594	0.438
PT NJ	0.020	0.037	0.060	0.100	0.034	0.023	0.073	0.054
Na	0.087	0.160	0.243	0.396	0.155	0.086	0.230	0.231
Sm	0.033	0.049	0.054	0.075	0.048	0.018	0.036	0.079
Eu	0.013	0.022	0.012	0.019	0.028	0.003	0.009	0.03/
Gđ	0.072	0.098	0.062	0.06/	0.05/	0.024	0.046	0.121
10	0.017	0.022	0.011	0.010	0.009	0.004	0.009	0.024
Dy	0.159	0.192	0.073	0.061	0.064	0.032	0.085	0.173
HO	0.042	0.051	0.016	0.014	0.015	0.008	0.022	0.035
Er	0.146	0.1/0	0.049	0.044	0.043	0.025	0.074	0.098
1m Vh	0.023	0.028	0.008	0.008	0.006	0.004	0.013	0.014
YD	0.1//	0.201	0.056	0.073	0.046	0.031	0.095	0.083
Lu M-#	0.028	0.031	0.009	0.014	0.008	0.005	0.015	0.011
₩Ig#	85.8	88.5	90.6	90.1	90.8	90.6	89.5	83.2

Hart & Zindler, 1986). Thus, peridotites of residual mantle origin should follow a progressive trend toward lower Al₂O₃/SiO₂ values (reflecting a higher degree of melt extraction) and higher MgO/SiO₂ values along a trend that parallels the terrestrial array (Fig. 6). The HSP rocks show a trend extending from the peridotites, which have more fertile $Al_2O_3/SiO_2 = 0.07-0.09$ and $MgO/SiO_2 = 1.1-1.14$, to more refractory peridotite $(Al_2O_3/SiO_2 = 0.00-0.05 \text{ and } MgO/SiO_2 = 0.75-1.1);$ they all plot along or above the terrestrial array. Their high MgO/SiO₂ values reflect their higher olivine/pyroxene ratios. Some samples (UH-28 and UH-30) record low Al_2O_3/SiO_2 values (~0.00), reflecting their more depleted (refractory) compositions; they are not plotted. On this diagram (Fig. 6), some samples show lower values of MgO/SiO₂ and Al₂O₃/SiO₂, indicating that they have experienced MgO loss or SiO_2 addition relative to normal mantle melting residues (Marchesi *et al.* 2013).

Most of these rocks have REE patterns that are subparallel to each other (Fig. 7a), although two samples exhibit REE patterns that are accompanied by slightly bulged shapes around their MREE segments. All of the peridotite samples are characterized by variously depleted REE patterns compared to the depleted MORB mantle (DMM), and their total REE contents range from 0.5 to 6.2 ppm. The presence of slightly positive Eu anomalies in some samples (Niu, Langmuir & Kinzler, 1997) is due to reactions with plagioclase-bearing rocks, such as oceanic gabbro (e.g. Klinkhammer *et al.* 1994; Douville *et al.* 2002; Chen *et al.* 2013). Recent experimental studies demonstrate that the chlorinity

Table 2. Conti	nued
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Sample Rock type Location	UH-21 Cpx-hz Sahneh	UH-22 Cpx-hz Sahneh	UH-23 Cpx-hz Sahneh	UH-24 Du Sahneh	UH-26 Cpx-hz Harsin	UH-28 Cpx-hz Harsin	UH-30 Cpx-hz Sahneh	UH-31 Cpx-hz Sahneh
$\overline{SiO_{2}}$ (wt%)	39.24	39.01	43.48	43 36	41.04	41 53	39.85	39.21
TiO ₂	0.01	0.01	0.01	0.01	0.01	0.02	0.02	0.06
AlaOa	2.87	0.01	0.01	0.01	0.01	0.02	0.02	0.06
Fe ₂ O ₂	9.84	10.84	8.09	9.08	10.06	10.30	10.93	10.85
MnO	0.14	0.15	0.18	0.13	0.15	0.14	0.13	0.14
MgO	40.02	42.52	35.48	13 23	13.87	14 20	42.58	13 25
CaO	3.00	-2.52	8 18	43.23	+5.87	0.96	42.58	-9.25
Na-O	0.01	0.82	0.04	0.93	0.10	0.90	0.49	0.50
K ₂ O	0.01	0.11	0.04	0.02	0.10	0.09	0.09	0.18
R ₂ O P.O.	0.00	0.00	0.01	0.00	0.00	0.00	0.00	0.00
$\Gamma_2 O_5$	0.000	1 10	0.007	0.000	0.004	0.61	0.005	0.010
NiO	0.082	0.37	0.50	0.37	0.00	0.01	0.35	0.33
	4.36	6.45	4.35	2.45	2.18	1.08	5.23	5.03
TOTAL	100.63	101 43	100.45	100.52	100.10	100.27	100.07	100.22
$\frac{101AL}{Sc(nnm)}$	6.23	8 02	15.7	9.60	100.10	9.67	7.11	10.22
V (ppin)	10.0	22.4	15.7	22.0	10.1	20.2	20.6	25.0
v Cr	10.9	52.4 7560	43.2	2800	42.0	39.2 4200	2780	33.0
C	119	108	2430 64.2	02.8	4320	4200	2780	01.4
C0 Ni	110	108	572	92.0	105	2760	2720	91.4
INI Cu	1140	2000	373	2350	2720	2700	2750	2030
Cu 7	150	5.05	2.20	0.30	21.3	15.0	2.74	24.1
Zn	38.9	40.5	40.1	33.9 0.557	45.7	44.8	55.0 0.705	34.1 1.75
Dh Dh	1.10	0.088	0.717	0.557	0.041	0.720	0.703	1.73
KU Su	20.4	0.724	1.24	0.021	0.466	0.290	0.225	18 6
51 7.	50.4	0.340	27.2	5.50	0.334	0.127	5.21	40.0
ZI Nh	0.085	0.107	0.099	0.071	0.320	0.098	0.980	5.20
ND Ca	n.a	0.039	0.004	0.004	0.085	0.039	0.075	0.373
CS Do	0.018	0.008	0.207	0.041	0.005	0.003	0.000	0.025
Da Dh	5.12	0.204	2.24	1.22	0.182	0.105	0.02	2.38
PD TL	0.4/1	0.088	0.051	0.131	0.096	0.072	0.031	0.155
In	0.019	0.015	0.014	0.005	0.015	0.006	0.023	0.033
U	0.005	0.006	0.004	0.002	0.005	0.004	0.000	0.015
HI	0.005	0.006	0.004	0.003	0.012	0.003	0.027	0.098
la	0.010	0.012	0.012	0.008	0.038	0.008	0.009	0.033
$\frac{\mathbf{I}}{\mathbf{L}_{\mathbf{a}}$ (mmm)	0.033	0.247	0.217	1.29	0.282	0.11/	0.191	1.33
La (ppm)	0.820	0.528	0.192	1.25	0.026	0.006	0.062	0.434
Ce D	1.20	0.811	0.227	2.21	0.050	0.012	0.134	0.969
Pr	0.133	0.077	0.028	0.282	0.006	0.001	0.017	0.127
Na	0.485	0.247	0.098	1.19	0.027	0.005	0.079	0.593
Sm	0.102	0.045	0.022	0.264	0.009	0.001	0.019	0.16/
Eu	0.056	0.007	0.003	0.084	0.006	0.001	0.008	0.084
Gd	0.118	0.043	0.031	0.295	0.019	0.004	0.026	0.218
Tb	0.019	0.006	0.005	0.044	0.005	0.001	0.004	0.037
Dy	0.128	0.043	0.042	0.274	0.041	0.013	0.033	0.251
HO	0.026	0.010	0.009	0.052	0.011	0.004	0.008	0.056
Er	0.076	0.032	0.026	0.124	0.043	0.020	0.028	0.160
Im	0.011	0.005	0.004	0.012	0.009	0.004	0.005	0.024
Yb	0.079	0.041	0.033	0.055	0.064	0.040	0.044	0.165
Lu	0.012	0.007	0.005	0.008	0.012	0.008	0.008	0.025
Mg#	89.2	88.6	89.7	90.4	89.6	89.5	88.5	88.8

and redox potential of the fluid is a major controlling factor on LREE and Eu mobility (Allen & Seyfried, 2005; Paulick *et al.* 2006; Deschamps *et al.* 2013) and that the presence of plagioclase is not required for the generation of LREE-enriched fluid compositions with positive Eu anomalies.

The abundances of incompatible trace elements on the primitive mantle-normalized diagram (Fig. 7b) show elevated abundances of U, Pb, Sr, Cs and Rb and small negative Ba and Zr anomalies compared to their neighbouring elements. Except for a few elements, such as Cs, Pb and Sr, in some samples, the trace element contents in the HSP samples are depleted compared to those of the primitive mantle (Sun & McDonough, 1989). This feature suggests the residual mantle origin of their protolith. Assessing the abundances of the incompatible trace elements normalized to the primitive mantle (Fig. 7b) reveals that some samples show a decoupling between Nb and Ta and are more enriched in Ta. This is consistent with depleted peridotites from abyssal and continental passive margins, whereas two samples show higher depletion on MREE and HREE compared to abyssal peridotites and the REE patterns reaching up to the field of typical supra-subduction peridotites (Parkinson & Pearce, 1998; Uysal *et al.* 2016).

7. Sr-Nd isotopic ratios

The ⁸⁷Sr/⁸⁶Sr and ¹⁴³Nd/¹⁴⁴Nd isotopic ratios of the HSP rocks are shown in Table 3. Their initial ratios are calculated based on an age of 80 Ma, which

Table 2. Continued

	Sample Rock type	UH-42 Cpx-hz	UH-46 Lhz	UH-53 Cpx-hz	SH1 Du	SH2 Du	SH3 Du	TB2 Du	UM1 Du
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Location	Sahneh	Sahneh	Harsin	Harsin	Harsin	Harsin	Harsin	Harsin
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	SiO ₂ (wt%)	41.35	42.62	41.07	40.1	41.1	40.6	40.5	40.0
$ \begin{array}{c c c c c c c c c c c c c c c c c c c $	TiO ₂	0.04	0.01	0.02	0.01	0.01	0.01	0.02	0.03
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Al_2O_3	1.37	3.70	1.87	0.37	1.08	0.62	1.14	1.15
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Fe ₂ O ₃	10.52	10.52	10.77	8.86	9.21	8.51	7.62	8.77
$ Mg0 45.17 35.17 41.33 43.5 45.9 45.5 45.9 44.3 \\ CaO 0.09 0.03 0.03 0.02 0.02 0.03 0.06 0.06 \\ K_O 0.00 0.00 0.00 0.01 0.01 0.01 0.02 0.01 \\ P_O, 0.011 0.008 0.009 0.010 0.010 0.010 0.01 0.02 0.01 \\ P_C, 0.011 0.008 0.009 0.010 0.010 0.010 0.01 0.01 \\ T_C, O_1 * * * * * * * * * $	MnO	0.10	0.11	0.18	0.11	0.12	0.12	0.08	0.12
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	MgO	45.17	35.17	41.33	43.5	45.9	45.5	45.9	44.3
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	CaO	0.29	1.01	0.26	0.62	1.29	0.72	0.23	1.12
	Na ₂ O	0.09	0.03	0.03	0.02	0.02	0.03	0.06	0.06
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	$K_2 \tilde{O}$	0.00	0.00	0.00	0.01	0.01	0.01	0.02	0.01
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	P_2O_5	0.011	0.008	0.009	0.010	0.010	0.010	nd	nd
$\begin{array}{c c c c c c c c c c c c c c c c c c c $	Cr_2O_3	*	*	*	0.39	0.32	0.45	0.36	0.39
$ \begin{array}{c c c c c c c c c c c c c c c c c c c $	NiO	*	*	*	*	*	*	*	*
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	LOI	1.24	6.83	4.59	6.82	1.29	3.34	4.20	3.82
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	TOTAL	100.18	100.01	100.13	100.82	100.36	99.92	100.11	99.76
V No 32.7 2.57 20.5 5 5 5 5 $n.d.$ $n.d.$ Cr 651 95.6 3630 2590 2200 3030 2490 2900 Co 13.9 15.6 126 102 119 111 96.5 110 Ni 112 161 1390 2150 2380 2260 2080 2110 Cu 41.1 25.6 3.35 5 8 5 5 13 Zn 10.2 11.0 72.4 43 52 47 40 52 Ga 12.0 12.5 2.89 0.5 1.1 0.8 16 1.4 Rb 0.854 0.105 0.893 0.2 0.2 0.3 0.4 0.9 Sr 202 178 3.54 6.7 2.9 0.9 9.2 5.6 Zr 1.52 0.359 1.28 $n.d$ $n.d$ $n.d$ 2.3 Nb 0.205 0.095 0.186 0.2 0.2 0.2 $n.d$ $n.d$ Cs 0.131 0.026 0.020 0.01 0.01 0.01 0.07 0.02 Ba 4.46 5.17 1.31 2.1 5.1 1.8 2.7 2.2 Pb 0.176 0.145 0.221 5 5 5 $n.d$ $n.d$ Th 0.230 0.76 0.051 0.05 0.05 0.05 0.06 0.05 <td< td=""><td>Sc (ppm)</td><td>*</td><td>*</td><td>*</td><td>*</td><td>*</td><td>*</td><td>*</td><td>*</td></td<>	Sc (ppm)	*	*	*	*	*	*	*	*
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	V	32.7	2.57	20.5	5	5	5	n.d.	n.d.
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Cr	651	95.6	3630	2590	2200	3030	2490	2900
Ni112161139021502380226020802110Cu41.125.63.35585513Zn10.211.072.44352474052Ga12.012.52.890.51.10.8161.4Rb0.8540.1050.8930.20.20.30.40.9Sr2021783.546.72.90.99.25.6Zr1.520.3591.28n.dn.dn.dn.dn.dCs0.1310.0260.0200.010.010.010.070.02Ba4.465.171.312.15.11.82.72.2Pb0.1760.1450.221555n.dn.dCu0.0180.0010.0070.050.050.050.060.05U0.0180.0010.0070.050.050.05n.dn.dTa0.1340.0310.018n.dn.dn.dn.dn.dY1.240.0990.2840.50.50.50.7n.dCe0.6270.2310.2540.50.50.50.9n.dPr0.0700.0340.0330.030.030.030.030.030.03Nd0.3450.1480.1450.10.10.1	Со	13.9	15.6	126	102	119	111	96.5	110
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Ni	112	161	1390	2150	2380	2260	2080	2110
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	Cu	41.1	25.6	3.35	5	8	5	5	13
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Zn	10.2	11.0	72.4	43	52	47	40	52
Rb 0.854 0.105 0.893 0.2 0.2 0.3 0.4 0.9 Sr 202 178 3.54 6.7 2.9 0.9 9.2 5.6 Zr 1.52 0.359 1.28 n.dn.dn.d 2 3 Nb 0.205 0.095 0.186 0.2 0.2 0.2 n.dn.dCs 0.131 0.026 0.020 0.01 0.01 0.01 0.07 0.02 Ba 4.46 5.17 1.31 2.1 5.1 1.8 2.7 2.2 Pb 0.176 0.145 0.221 5 5 $n.d$ $n.d$ Th 0.230 0.176 0.051 0.05 0.05 0.05 0.06 0.05 U 0.018 0.001 0.007 0.05 0.05 0.05 $n.d$ $n.d$ Ta 0.134 0.031 0.018 $n.d$ $n.d$ $n.d$ $n.d$ $n.d$ Y 1.24 0.099 0.289 0.5 0.5 0.5 0.7 $n.d$ Ce 0.627 0.231 0.254 0.5 0.5 0.5 0.9 $n.d$ Nd 0.345 0.148 0.145 0.1 0.1 0.1 0.7 0.2 Sm 0.126 0.29 0.36 0.03 0.03 0.03 0.15 $n.d$ Md 0.345 0.148 0.145 0.1 0.1 0.1 0.1 0.16 <	Ga	12.0	12.5	2.89	0.5	1.1	0.8	16	1.4
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Rb	0.854	0.105	0.893	0.2	0.2	0.3	0.4	0.9
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Sr	202	178	3.54	6.7	2.9	0.9	9.2	5.6
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Zr	1.52	0.359	1.28	n.d	n.d	n.d	2	3
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	Nb	0.205	0.095	0.186	0.2	0.2	0.2	n.d	n.d
Ba4.465.171.312.15.11.82.72.2Pb0.1760.1450.221555n.dn.dTh0.2300.1760.0510.050.050.050.060.05U0.0180.0010.0070.050.050.05n.dn.dHf0.0520.0030.0360.20.20.2n.dn.dTa0.1340.0310.018n.dn.dn.dn.dn.dY1.240.0990.2890.50.50.50.7n.dLa (ppm)0.2560.1520.1300.50.50.50.9n.dCe0.6270.2310.2540.50.50.50.9n.dNd0.3450.1480.1450.10.10.10.70.28Sm0.1260.0290.0360.030.030.030.030.1n.dGd0.3450.1480.1450.10.10.10.70.28Sm0.1260.0290.0360.030.030.030.030.030.03Gd0.2000.0300.0430.050.050.050.09n.dMd0.360.0040.010.010.01n.dn.dMd0.0360.0040.050.050.050.050.09n.dDy0.2470.0210.047	Cs	0.131	0.026	0.020	0.01	0.01	0.01	0.07	0.02
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Ba	4.46	5.17	1.31	2.1	5.1	1.8	2.7	2.2
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Pb	0.176	0.145	0.221	5	5	5	n.d	n.d
U0.0180.0010.0070.050.050.05n.dn.dHf0.0520.0030.0360.20.20.2n.dn.dTa0.1340.0310.018n.dn.dn.dn.dn.dY1.240.0990.2890.50.50.50.50.80.1La (ppm)0.2560.1520.1300.50.50.50.7n.dCe0.6270.2310.2540.50.50.50.9n.dNd0.3450.1480.1450.10.10.10.70.2Sm0.1260.0290.0360.030.030.030.030.15n.dGd0.2000.0300.0430.050.050.050.09n.dGd0.2000.0040.0080.010.01n.dn.dn.dDy0.2470.0210.0470.050.050.050.10.07Ho0.0500.0040.0110.010.010.010.01n.dTh0.0200.0010.0210.030.030.030.030.050.06Th0.0200.0010.0040.010.010.010.010.010.01Ma0.0200.0010.0040.010.010.010.010.010.01Mg#88.686.988.490.790.891.492.3 </td <td>Th</td> <td>0.230</td> <td>0.176</td> <td>0.051</td> <td>0.05</td> <td>0.05</td> <td>0.05</td> <td>0.06</td> <td>0.05</td>	Th	0.230	0.176	0.051	0.05	0.05	0.05	0.06	0.05
Hf 0.052 0.003 0.036 0.2 0.2 0.2 $n.d$ $n.d$ $n.d$ Ta 0.134 0.031 0.018 $n.d$ $n.d$ $n.d$ $n.d$ $n.d$ $n.d$ $n.d$ Y 1.24 0.099 0.289 0.5 0.5 0.5 0.5 0.8 0.1 La (ppm) 0.256 0.152 0.130 0.5 0.5 0.5 0.7 $n.d$ Ce 0.627 0.231 0.254 0.5 0.5 0.5 0.9 $n.d$ Pr 0.070 0.034 0.033 0.03 0.03 0.03 0.15 $n.d$ Nd 0.345 0.148 0.145 0.1 0.1 0.1 0.7 0.2 Sm 0.126 0.029 0.036 0.03 0.03 0.03 0.1 $n.d$ Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 $n.d$ Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.011 0.01 0.01 0.01 0.01 0.01 0.01 Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.11 0.07 Ho 0.050 0.004 0.011 0.01 0.01 0.01 0.01 0.01 0.01 Hu 0.007 0.020 0.05 0.066 0.04 0.11 0.11	U	0.018	0.001	0.007	0.05	0.05	0.05	n.d	n.d
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Hf	0.052	0.003	0.036	0.2	0.2	0.2	n.d	n.d
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Та	0.134	0.031	0.018	n.d	n.d	n.d	n.d	n.d
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	Y	1.24	0.099	0.289	0.5	0.5	0.5	0.8	0.1
Ce 0.627 0.231 0.254 0.5 0.5 0.5 0.9 n.d Pr 0.070 0.034 0.033 0.03 0.03 0.03 0.15 n.d Nd 0.345 0.148 0.145 0.1 0.1 0.1 0.7 0.2 Sm 0.126 0.029 0.036 0.03 0.03 0.03 0.1 n.d 0.1 Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 n.d Tb 0.036 0.043 0.05 0.05 0.05 0.09 n.d Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.01 0.01 0.02 0.02 Er 0.141 0.009 0.021 0.03 0.03 0.03 0.05 0.06 Tm 0.020 0.001 0.01 0.01 0.01	La (ppm)	0.256	0.152	0.130	0.5	0.5	0.5	0.7	n.d
Pr 0.070 0.034 0.033 0.03 0.03 0.03 0.15 n.d Nd 0.345 0.148 0.145 0.1 0.1 0.1 0.7 0.2 Sm 0.126 0.029 0.036 0.03 0.03 0.03 n.d 0.08 Eu 0.129 0.112 0.011 0.03 0.03 0.03 0.1 n.d Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 n.d Tb 0.036 0.004 0.008 0.01 0.01 n.d n.d Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.01 0.01 0.02 0.02 Er 0.141 0.009 0.021 0.03 0.03 0.05 0.06 Tm 0.020 0.001 0.01 0.01 0.01 0.01 n.d	Ce	0.627	0.231	0.254	0.5	0.5	0.5	0.9	n.d
Nd 0.345 0.148 0.145 0.1 0.1 0.1 0.7 0.2 Sm 0.126 0.029 0.036 0.03 0.03 0.03 n.d 0.08 Eu 0.129 0.112 0.011 0.03 0.03 0.03 0.1 n.d Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 n.d Tb 0.036 0.004 0.008 0.01 0.01 n.d n.d Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.011 0.01 0.02 0.02 0.02 Er 0.141 0.009 0.21 0.03 0.03 0.03 0.05 0.06 Tm 0.020 0.001 0.001 0.01 0.01 n.d n.d Yb 0.119 0.007 0.020 0.05 0.06 0.04	Pr	0.070	0.034	0.033	0.03	0.03	0.03	0.15	n.d
Sm 0.126 0.029 0.036 0.03 0.03 0.03 n.d 0.08 Eu 0.129 0.112 0.011 0.03 0.03 0.03 0.1 n.d Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 n.d Tb 0.036 0.004 0.008 0.01 0.01 0.01 n.d n.d Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.01 0.01 0.02 0.02 0.02 Er 0.141 0.009 0.021 0.03 0.03 0.03 0.05 0.06 Tm 0.020 0.001 0.001 0.01 0.01 n.d1 0.11 0.11 Lu 0.017 0.001 0.004 0.01 0.01 0.01 0.01 n.d1 Lu 0.017 0.001 0.004 0.	Nd	0.345	0.148	0.145	0.1	0.1	0.1	0.7	0.2
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	Sm	0.126	0.029	0.036	0.03	0.03	0.03	n.d	0.08
Gd 0.200 0.030 0.043 0.05 0.05 0.05 0.09 n.d Tb 0.036 0.004 0.008 0.01 0.01 0.01 n.d n.d Dy 0.247 0.021 0.047 0.05 0.05 0.05 0.1 0.07 Ho 0.050 0.004 0.011 0.01 0.02 0.01 0.02 0.02 Er 0.141 0.009 0.021 0.03 0.03 0.03 0.05 0.06 Tm 0.020 0.001 0.01 0.01 0.01 0.01 n.d Vb 0.119 0.007 0.020 0.05 0.06 0.04 0.11 0.11 Lu 0.017 0.001 0.004 0.01 0.01 0.01 0.02 Mg# 88.6 86.9 88.4 90.7 90.8 91.4 92.3 90.9	Eu	0.129	0.112	0.011	0.03	0.03	0.03	0.1	n.d
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	Gd	0.200	0.030	0.043	0.05	0.05	0.05	0.09	n.d
Dy0.2470.0210.0470.050.050.050.10.07Ho0.0500.0040.0110.010.020.010.020.02Er0.1410.0090.0210.030.030.030.050.06Tm0.0200.0010.0040.010.010.010.01n.dYb0.1190.0070.0200.050.060.040.110.11Lu0.0170.0010.0040.010.010.010.02Mg#88.686.988.490.790.891.492.390.9	Tb	0.036	0.004	0.008	0.01	0.01	0.01	n.d	n.d
Ho0.0500.0040.0110.010.020.010.020.02Er0.1410.0090.0210.030.030.030.050.06Tm0.0200.0010.0040.010.010.010.010.01n.dYb0.1190.0070.0200.050.060.040.110.11Lu0.0170.0010.0040.010.010.010.02Mg#88.686.988.490.790.891.492.390.9	Dy	0.247	0.021	0.047	0.05	0.05	0.05	0.1	0.07
Er0.1410.0090.0210.030.030.030.050.06Tm0.0200.0010.0040.010.010.010.01n.dYb0.1190.0070.0200.050.060.040.110.11Lu0.0170.0010.0040.010.010.010.010.02Mg#88.686.988.490.790.891.492.390.9	Ho	0.050	0.004	0.011	0.01	0.02	0.01	0.02	0.02
Tm0.0200.0010.0040.010.010.010.01n.dYb0.1190.0070.0200.050.060.040.110.11Lu0.0170.0010.0040.010.010.010.010.02Mg#88.686.988.490.790.891.492.390.9	Er	0.141	0.009	0.021	0.03	0.03	0.03	0.05	0.06
Yb 0.119 0.007 0.020 0.05 0.06 0.04 0.11 0.11 Lu 0.017 0.001 0.004 0.01 0.01 0.01 0.01 Mg# 88.6 86.9 88.4 90.7 90.8 91.4 92.3 90.9	Tm	0.020	0.001	0.004	0.01	0.01	0.01	0.01	n.d
Lu0.0170.0010.0040.010.010.010.010.01Mg#88.686.988.490.790.891.492.390.9	Yb	0.119	0.007	0.020	0.05	0.06	0.04	0.11	0.11
Mg# 88.6 86.9 88.4 90.7 90.8 91.4 92.3 90.9	Lu	0.017	0.001	0.004	0.01	0.01	0.01	0.01	0.02
	Mg#	88.6	86.9	88.4	90.7	90.8	91.4	92.3	90.9

*=not measured; .d.=not detected.

Aziz, Elias & Aswad (2011) reported from neighbouring Penjween serpentinized peridotites using the Rb– Sr method (80-100 Ma). Braud (1978) and Shahidi & Nazari (1997) have suggested a Jurassic to Late Cretaceous age for the Harsin–Sahneh ophiolite complex based on its tectonic and stratigraphic relationships. Their ⁸⁷Sr/⁸⁶Sr_(i=initial) and ¹⁴³Nd/¹⁴⁴Nd_(i) ratios range from 0.7038 to 0.7109 and from 0.51214 to 0.51285, respectively. Their ϵ Nd_(t) values range from –7.5 to +7.8 (Table 3), thus reflecting the contribution of continental material to the source of the HSP rocks. The ⁸⁷Sr/⁸⁶Sr_(i) ratios of the HSP rocks are significantly more radiogenic compared to MORB (0.7025), whereas the ¹⁴³Nd/¹⁴⁴Nd_(i) ratios of the HSP rocks are lower than that of MORB (0.51315). This suggests that the HSP rocks originated from a depleted mantle source that was affected by alteration reactions (as indicated by their high 87 Sr/ 86 Sr_(i) ratios).

On the ⁸⁷Sr/⁸⁶Sr_(i) versus ε Nd_(t) diagram, the samples plot near the mantle array, and they extend from the MORB (mid-oceanic ridge) field toward the EMII (continental sediments associated with a subducting slab) field and a seawater component (Fig. 8a, b, c). These diagrams show that the HSP rocks were not strongly affected by seawater alteration and the HSP rocks were altered by seawater along deep faults and fractures in an extensional basin. In Figure 8a, the HSP rocks exhibit minor overlapping with the Harsin ophiolite domain due to their highly radiogenic Sr isotopic compositions and less radiogenic Nd isotopic

W/R(Closed- W/R(Open- system) system)	2.27 1.05 1.37 0.78 2.24 1.05 1.50 0.83	0.20 0.20 0.20 0.79 0.79 0.54 0.88 0.54 0.58 0.54 0.58 3.15 1.35 1.35 1.36 0.70 0.70	2.70 6.98 0.95 0.12 0.12 0.11 0.02 0.74 0.52 0.52 1.05 0.66
f(Rb/ Sr)	3.26 2.61 4.63 4.88	2.20 2.24 18.8 4.18 -0.23 72.2 0.59 0.59	$\begin{array}{c} 11.7\\ 27.3\\ 27.3\\ 1.43\\ -0.54\\ -0.98\\ -0.98\\ 7.83\\ 7.83\\ 4.6\\ 4.6\end{array}$
f(Sm/ Nd)	0.16 - 0.06 - 0.32 - 0.42	-0.00 -0.52 -0.52 -0.52 -0.35 -0.3	$\begin{array}{c} 0.07\\ -0.12\\ -0.26\\ -0.13\\ 0.13\\ -0.39\\ -0.24\\ -0.08\\ -0.23\end{array}$
εNd (t)	$\begin{array}{c} -1.8\\ 1.4\\ 2.8\end{array}$	-2.8 -0.6 -3.9 -6.6 -1.0	6.3 6.6 6.6 3.7 7.8 7.8 5.7 - 7.5 - 7.5
$^{143}_{144}Nd/$	0.51243 0.51260 0.51267	0.51266 0.51249 0.51272 0.51219 0.51219 0.51247	0.51284 0.51286 0.51238 0.51271 0.51271 0.51271 0.51281 0.51276 0.51214
$\pm 1SE$	$\begin{array}{c} 0.000045\\ 0.000008\\ 0.000006\end{array}$	$\begin{array}{c} 0.000023\\ 0.000036\\ 0.000009\\ 0.000013\\ 0.000013\\ 0.000062 \end{array}$	$\begin{array}{c} 0.000004\\ 0.000005\\ 0.000004\\ 0.000004\\ 0.000008\\ 0.000029\\ 0.000018\\ 0.000018\\ 0.000018\end{array}$
$^{143}_{144}Nd/$	0.512536 0.512680 0.512729	0.512721 0.512534 0.512817 0.512244 0.512524	0.512939 0.512938 0.512451 0.512788 0.512788 0.512788 0.512868 0.512826 0.512826 0.512222
¹⁴⁷ Sm/ ¹⁴⁴ Nd	0.228 0.185 0.134 0.115 0.115	0.130 0.126 0.206 0.128 0.111 0.134	0.210 0.173 0.146 0.170 0.170 0.221 0.120 0.149 0.181 0.242
Sm	0.033 0.049 0.054 0.075 0.075	0.048 0.018 0.036 0.079 0.045 0.022 0.022	$\begin{array}{c} 0.009\\ 0.001\\ 0.019\\ 0.167\\ 0.126\\ 0.029\\ 0.036\\ 0.030\\ 0.080\end{array}$
ΡN	0.087 0.160 0.243 0.396	0.130 0.086 0.231 0.231 0.231 0.247 0.098	$\begin{array}{c} 0.027\\ 0.005\\ 0.079\\ 0.593\\ 0.345\\ 0.148\\ 0.146\\ 0.100\\ 0.200\\ 0.200 \end{array}$
$^{87}_{\rm Sr/}{ m Sr}_{ m (i)}$	0.70750 0.70682 0.70747 0.70693	0.70604 0.70604 0.70388 0.70595 0.70595 0.70614 0.70730 0.70805 0.70682	0.7056 0.70829 0.71091 0.70827 0.70358 0.70358 0.70580 0.70580 0.70640
± 1SE	0.000008 0.000006 0.000008 0.000008	0.000007 0.000006 0.000006 0.000006 0.000006 0.000006	$\begin{array}{c} 0.000013\\ 0.000017\\ 0.000006\\ 0.000006\\ 0.000006\\ 0.000006\\ 0.000005\\ 0.000005\\ 0.000005\\ 0.000006\end{array}$
$^{87}{ m Sr}/_{ m 8^6}{ m Sr}_{ m (p)}$	0.707533 0.706848 0.707518 0.706975	0.704422 0.706068 0.706068 0.705991 0.706151 0.707899 0.708067	0.707734 0.708523 0.710931 0.706272 0.705272 0.705875 0.705875 0.705875 0.7058443
⁸⁷ Rb/ ⁸⁶ Sr	0.353 0.299 0.466 0.486	0.254 0.268 0.429 0.064 0.132 0.327	$\begin{array}{c} 1.05\\ 2.34\\ 0.201\\ 0.038\\ 0.012\\ 0.730\\ 0.465\\ 0.465\end{array}$
Sr	4.59 6.49 5.67 10.8 4.18	4.16 4.96 7.90 30.4 0.346 5.0 5.50	$\begin{array}{c} 0.816\\ 0.115\\ 3.21\\ 48.6\\ 202\\ 178\\ 3.54\\ 0.86\\ 5.60\end{array}$
Rb	0.559 0.670 0.912 1.82	0.423 0.460 1.17 0.670 0.724 1.24 1.24	$\begin{array}{c} 0.296\\ 0.093\\ 0.634\\ 0.854\\ 0.105\\ 0.105\\ 0.140\\ 0.140\\ 0.900 \end{array}$
Sample	UH-3 UH-5 UH-8 UH-9	UH-17 UH-17 UH-20 UH-21 UH-22 UH-22	UH-26* UH-28* UH-30* UH-31* UH-31* UH-42 UH-42 UH-45 SH1* UM1

Table 3. Rb-Sr and Sm-Nd isotopic compositions of the HSP rocks



Figure 6. (Colour online) Plot of the HSP rocks on Al_2O_3/SiO_2 versus MgO/SiO₂ diagram (Jagoutz *et al.* 1979; Hart & Zindler, 1986). The HSP rocks extend from peridotites in more fertile $Al_2O_3/SiO_2 = 0.07-0.09$ and MgO/SiO₂ = 1.1-1.14 and more refractory peridotites ($Al_2O_3/SiO_2 = 0.00-0.05$ and MgO/SiO₂ = 0.75-1.1); they all plot along or above the terrestrial array.

compositions compared to other domains. Figure 8d shows the relationship between the Nd isotopic compositions and (Ce/Yb)_N ratios of the HSP rocks. In this diagram, the DP trend is characterized by higher ϵ Nd values (ϵ Nd > +10) and more depleted LREE patterns ((Ce/Yb)_N <0.5) compared to the MORB source ($\epsilon Nd = +10$ (White & Hofmann, 1982) and $(Ce/Yb)_N = 0.66$ (Sun & McDonough, 1989)). The HSP samples exhibit enriched REE signatures, and they plot in the EP field. The EP field is formed by mixing between an enriched component and residual peridotites (McCulloch & Chappel, 1982); only continental crustal material shows such enriched characteristics (i.e. fluids or melts derived from the continental crust have an average $(Ce/Yb)_N = 4$). Thus, the minor mantle metasomatism of the HSP rocks was derived from continental crustal materials associated with the subducted slab (Fig. 8b, c, d). Additionally, the calculated values of the Sm/Nd and Rb/Sr differentiation factors ($f_{(Sm/Nd)}$ and $f_{(Rb/Sr)}$ (McCulloch & Wasserburg, 1978)) are shown in Table 3. The careful use of these equations ($f_{(Sm/Nd)}$ and $f_{(Rb/Sr)}$ values) allows us to identify the mixing of a source with continental crustal material. The $f_{(Sm/Nd)}$ values range from -0.52 to 0.23, and the $f_{(Rb/Sr)}$ values range from -1 to 72, which confirms that the mantle source of the HSP rocks experienced the minor contribution of continental crustal materials associated with the subducting slab.

8. Discussion

8.a. Evidence of melt depletion

The petrographic, major and trace element characteristics of the HSP rocks indicate that they represent the refractory and depleted mantle residues following partial melting. The geochemical and petro-



Figure 7. (Colour online) (a) Chondrite-normalized abundances of REE in the HSP rocks (Sun & McDonough, 1989). The supra-subduction zone (SSZ) peridotite field is from Parkinson & Pearce (1998). The DMM field is from Sun & McDonough (1989). Fields of abyssal peridotites and melt/rock interaction are from Deschamps et al. (2013). The solid line field (abyssal peridotites + melt/rock interactions) shows characteristics of refertilized protolith after melt/rock interactions. (b) Trace element abundances normalized by primitive mantle values (Sun & McDonough, 1989). The solid line field (abyssal peridotites + melt/rock interactions) shows characteristics of refertilized protolith after melt/rock interactions. The field of passive margin serpentinized peridotites is based on data from Newfoundland and the Iberian abyssal plain from Seifert & Brunotte (1996), Niu (2004), Savov et al. (2005, 2007), Deschamps et al. (2011, 2012) and Kodolányi et al. (2012).

graphical evidence includes: (1) the orthopyroxene grains with exsolution lamellae of clinopyroxene and spinel or between orthopyroxene and olivine (Fig. 4h, i), olivines with kink band features and partial replacement of orthopyroxene by olivine (Fig. 4b); (2) their low ¹⁴³Nd/¹⁴⁴Nd_(i) ratios (0.51208–0.51285); (3) the forsteritic composition of primary olivine (Fo: 91.1–94.9); (4) the depletion of the HSP samples in terms of their major oxide (Al₂O₃, TiO₂, CaO and MnO),



Figure 8. (Colour online) (a) The variations in 87 Sr/ 86 Sr_(i) versus ϵ Nd_(t). The samples plot near the mantle array and extend from the field of the depleted mantle toward the fields of the EMII end-member component and seawater component. (b) The variations in 87 Sr/ 86 Sr_(i) versus 143 Nd/ 144 Nd_(i). The field of mantle components is based on data compiled by Stracke, Bizimis & Salters, (2003) for Atlantic MORB (DMM). Saint Helena (HIMU), Samoa and Society (EMII) and Pitcairn (EMI). The samples extend toward the EMII end-member, suggesting a minor contribution from EMII to the depleted mantle in their sources. The Sr and Nd isotopic compositions of seawater are from the literature: Palmer & Edmond (1989) for Sr and Nd compositions and Piepgras & Wasserburg (1987) and Tachikawa, Jeandel & Roy Barman, (1999) for Nd isotopic compositions. The fields of the Eocene Kamyaran ophiolite and Late Cretaceous Mawat ophiolite are from Azizi *et al.* (2011b) and Azizi *et al.* (2013), the data for the Late Cretaceous Neyriz ophiolite are from Shafaii Moghadam *et al.* (2014), the data for the Late Cretaceous Balvard–Baft ophiolite are from Shafaii Moghadam *et al.* (2013), and the data for the Harsin ophiolite are from Nouri (2016). (c) Sr isotope versus Ba/Rb diagram showing mixing between upper/lower crust and upper mantle peridotites. The Ba/Rb ratio of the lower crust is based on mafic granulites from the Dabie orogeny in Hou (2003). The Ba and Rb data for the UM and upper crust are from Sun & McDonough (1989). The Sr–Nd data for UM (upper mantle peridotite) are from Jahn *et al.* (1999), and the Sr–Nd data for the lower and upper crust are from Xu *et al.* (2004) and Jahn *et al.* (1999). (d) (Ce/Yb)_N versus ϵ Nd. Star indicates the (Ce/Yb)_N and ϵ Nd compositions of the present-day MORB (White & Hofmann, 1982; Sun & McDonough, 1989).

incompatible trace element (V, Sc, Y, Zn) and HREE contents; and (5) their variously depleted REE patterns compared to the depleted mantle on chondriteand primitive mantle-normalized diagrams. Additionally, their observed whole-rock geochemical compositions are consistent with those of refractory and depleted mantle residues, which can also be seen on the MgO/SiO2 versus Al2O3/SiO2 diagram, where these samples lie on the depleted end of the terrestrial melting trend (Fig. 6). On the TiO_2 versus Cr_2O_3 diagram (Fig. 5a), the spinel compositions of the HSP rocks fall within the field of depleted oceanic peridotites (tectonites). This can also be seen on the TiO₂ versus Cr# diagram for spinel (Fig. 9), where the HSP spinels fall within the field of refractory abyssal peridotites defined by Kelemen et al. (1997) and Suhr et al. (2003); additionally, indicating that they represent mantle residues.

8.b. Water-rock interaction or melt refertilization?

The chemical compositions of the HSP rocks exhibit similar properties to refractory peridotites. Generally, the refractory peridotites are residues of high degrees of partial melting; their whole-rock REE patterns should be characterized by strong LREE depletion and flat to enriched MREE–HREE contents (Johnson, Dick & Shimizu, 1990; Johnson & Dick, 1992). Nevertheless, the bulk compositions of the HSP rocks exhibit different patterns, with light rare earth elements (LREE) and large-ion lithophile elements (LILE) enrichment (Fig. 7). Therefore, there must be another



Figure 9. (Colour online) (a) TiO_2 versus Cr# diagram for the HSP spinels. The spinel samples fall within the field of refractory abyssal peridotite. The fields of refractory abyssal peridotites and Atlantic MORB are after Kelemen *et al.* (1997). (b) TiO_2 versus Cr# of spinels, the area for highly depleted peridotite from the Bay of Island ophiolite (Suhr *et al.* 2003), Josephine ophiolite (Morgan, Liang & Kelemen, 2008) and Iwanaidake peridotite (Kubo, 2002). MORB, island arc and boninite fields from Arai (1992), Kelemen, Shimizu & Salters (1995) and Dick & Natland (1996). Small black circles: dunites and impregnated peridotites from the Moho transition zone (MTZ) of the Oman ophiolite (Koga, Kelemen & Shimizu, 2001).

process, such as water–rock interaction (Paulick *et al.* 2006; Delacour *et al.* 2008; Chen *et al.* 2013) or melt refertilization (Niu, 1997, 2004; Kelemen, Kikawa & Miller, 2007; Chen *et al.* 2013), that can explain the LREE and LILE enrichments in the HSP rocks; these are discussed below.

8.b.1. Water-rock interaction ratios

In many studies, Sr isotopes have been demonstrated to be an important geochemical tool in the determination of fluid-rock interactions (Hart, Erlank & Kable, 1974; Jacobsen & Wasserburg, 1979; Hess, Bender & Schilling, 1991; Kempton, Hawkesworth & Fowler, 1991; Gillis, Coogan & Pedersen, 2005). A significant shift in Sr isotope compositions provides a constraint on the degree of seawater-rock interaction and fluid fluxes during the formation and alteration of these rocks. Therefore, when considering the interactions between fluids derived from seawater and the oceanic lithosphere, the ⁸⁷Sr/⁸⁶Sr ratio is an effective parameter for assessing water–rock (W/R) ratios (Hart, Erlank & Kable, 1974; Hess, Bender & Schilling, 1991; Kempton & Hunter, 1997; Delacour *et al.* 2008) and reflecting the extent of seawater–rock interaction in the oceanic lithosphere. It is necessary to calculate W/R ratios to determine the effect of hydrothermal alteration on the REE patterns of the HSP samples and to provide evidence for the circulation of seawater into the oceanic crust (Paulick *et al.* 2006; Delacour *et al.* 2008; Chen *et al.* 2013):

$$W/R_{\text{closedsystem}} = [(\varepsilon_{r}^{i} - \varepsilon_{r}^{f})/(\varepsilon_{r}^{f} - \varepsilon_{w}^{i})] \times (X_{r}/X_{w})$$
(1)

$$W/R_{\text{opensystem}} = (X_{\text{r}}/X_{\text{w}})$$
$$\times \ln[(\epsilon_{\text{r}}^{\text{ i}} - \epsilon_{\text{r}}^{\text{ f}})/(\epsilon_{\text{r}}^{\text{ f}} - \epsilon_{\text{w}}^{\text{ i}}) + 1] \qquad (2)$$

In this equation, *W* is the mass of the hydrothermal fluid, *R* is the mass of the rock being altered by the fluid, ε_r^{i} is the initial isotopic ratio of the rock (i.e. 0.702689 for mantle peridotite) (Workman & Hart, 2005), ε_r^{f} is the final isotopic ratio of the hydrothermally altered rock (Table 3), ε_w^{i} is the initial isotopic composition of hydrothermal fluid (seawater) (i.e. 0.70916) (Palmer & Edmond, 1989), X_r is the elemental composition of the mantle peridotite (i.e. 6.1 ppm) (Workman & Hart, 2005) and X_w is the elemental composition of the hydrothermal fluid (i.e. 8 ppm) (Palmer & Edmond, 1989).

When the water-rock ratio is greater than 100, incompatible elements become mobile and reflect the composition of the fluid phase rather than that of the parent rock (Bau, 1991; Gillis, Ludden & Smith, 1992; Delacour *et al.* 2008; Li *et al.* 2010). The *W*/*R* ratios of the HSP samples in both closed and open systems (Table 3) are low, ranging from 0.02 to 6.98 for closed systems and from 0.11 to 1.77 for open systems (Table 3). These data confirm that the minor circulation of seawater occurred along deep faults and fractures in an extensional basin. The *W*/*R* ratios of these rocks indicate that most REEs were unaffected by the hydrothermal alteration described in Section 7.

8.b.2. Melt refertilization

Melt refertilization is a common process in peridotites that has been studied in marine samples and ophiolitic successions (Godard, Bodinier & Vasseur, 1995; Seyler & Bonatti, 1997; Piccardo *et al.* 2003; Augustin *et al.* 2012). Trace element data from the serpentinized peridotites indicate that the geochemical characteristics of melt refertilization can even be preserved in altered ultramafic rocks (Paulick *et al.* 2006; Augustin *et al.* 2012; Chen *et al.* 2013; Manuella *et al.* 2016). The geochemical features related to melt refertilization include concordant increases in the concentrations of REEs and high field strength elements in bulk-rock peridotite samples (Niu, 2004; Augustin et al. 2012; Chen et al. 2013). This is due to the fact that REEs and HFSEs have similar solubilities in mafic melts, and this signature is transferred to peridotite. The trace element compositions of the serpentinized HSP rocks show positive trends on the Ce versus Th and La versus Nb diagrams (Fig, 10a, b), which indicates that melt-rock interaction is the dominant factor controlling these compositional variations, although some samples record evidence of minor fluid-rock interactions. These samples follow the trend of the global dataset, in which meltrock interaction is the dominant process. This relationship indicates that the variability in LREEs is due to melt-rock interaction with minor hydrothermal alteration, which affects some HFSEs. Moreover, the geochemical characteristics of the HSP rocks in the Th versus Gd/Lu diagram display a similar trend of meltrock interaction and exhibit enrichments in LREEs and HFSEs (Fig. 10c). In addition, these enrichments are confirmed by the La/Sm versus Sm/Yb diagram, which indicates that these samples have similar properties to the refertilized portion and exhibit enrichments in LREEs (Fig. 10d). Although Pb was previously considered to be a highly mobile element during hydrothermal alteration and dehydration in subduction zones (Niu & OHara, 2003), recent studies by Niu (2004) and Godard et al. (2008) of altered abyssal peridotites demonstrated that the geochemical behaviour of Pb during alteration and weathering was consistent with that of highly incompatible elements. Thus, the variations in the Nb versus Pb contents (Fig. 10e) are indicative of melt/fluid – rock interaction. In this diagram (Fig. 10e), the samples record evidence of melt refertilization; this trend is magmatic and indicates that the positive correlation of Nb and Pb is due to melt refertilization (Niu, 2004; Paulick et al. 2006). The scatter of some HSP samples on the Nb versus Pb diagram appears to be due to hydrothermal alteration modifying the Pb concentrations (Fig. 10e).

Moreover, most of the HSP rocks have similar contents of REE and incompatible elements to refertilized abyssal serpentinized peridotites (melt-rock interaction (Fig. 7)). They are enriched in LILEs and LREEs and show a decoupling between Nb and Ta, with more enrichment in Ta. In Figure 11, the REE in the HSP rocks have similar patterns to abyssal serpentinized peridotites that underwent melt-rock (MRI) and fluid-rock (FRI) interactions; however, they do not have similar patterns to the average hydrothermal sediments from the Mid-Atlantic Ridge (MAR) (Dias & Barriga, 2006). These processes are supported by the significant correlation between LREEs and HFSEs (Fig. 10), which reflects melt refertilization with minor fluid-rock interaction. Furthermore, the REE patterns of the HSP rocks on the chondrite-normalized diagram (Sun & McDonough, 1989) are almost flat and did not experience seawater alteration, as they exhibit negative Ce anomaly (Fig. 11). Therefore, we conclude that the enrichments in LREEs and other trace elements are clear evidence of melt–rock interaction (i.e. positive correlations between some LREEs and HFSEs), with some samples recording evidence of minor fluid–peridotite interaction.

8.c. Textures produced by melting and melt refertilization

Based on the above discussions, the textures and geochemistry of the HSP rocks offer evidences of cooling and decompression, partial melting and meltrock reactions; although each process produces a specific mineral texture, some textures are not limited to a specific process and instead reveal the collective effect of several processes. They typically consist of orthopyroxene grains with exsolution lamellae of clinopyroxene (Fig. 4c) and spinel, vermicular spinel located within orthopyroxene grains or between orthopyroxene and olivine porphyroclasts (Fig. 4h, i), and embayed and flame-like spinel grains (Fig. 4d, i) that are produced during cooling and decompression or by the partial melting of abyssal peridotites at low pressures (Edwards & Malpas, 1996; Pearce et al. 2000; Miller et al. 2003; Liu et al. 2010; Li et al. 2015). They are typical of oceanic mantle tectonites (Pearce et al. 2000; Tamura et al. 2008). Moseley (1984) and Tamura & Arai (2005) demonstrated that these types of cooling textures are widespread in both abyssal and SSZ peridotites, and are caused by a decrease in solubility of clinopyroxene and spinel in orthopyroxene due to cooling and decompression. In this stage, the partial melting of orthopyroxene leads to spinel/clinopyroxene intergrowths and the erosion of pyroxene grains (Fig. 4c) (Pearce et al. 2000; Miller et al. 2003; Liu et al. 2010). As partial melting continues and more melt is produced, vermicular, chain and flame-like spinel grains (Fig. 4h, i) are formed within orthopyroxene or olivine crystals. Moreover, high degrees of partial melting and melt-rock reactions can cause the breakdown of pyroxene and its pseudomorphic replacement of spinel. These evidences are widely observed in many peridotites and are related to cooling and decompression of mantle peridotites; they were produced by partial melting of primary or evolved mantle or melt-rock interaction (Siena & Coltorti, 1993; Edwards & Malpas, 1996; Roeder, Poustovetov & Oskarsson, 2001; Li et al. 2015).

Additionally, the vermicular grains of spinel located along grain boundaries and the interstitial clinopyroxene grains located between grains of olivine and orthopyroxene (Fig. 3f, g) are interpreted to represent the crystallization products of metasomatic melts (Morishita *et al.* 2007) and are also attributed to melt–rock interaction with incompletely extracted melts in abyssal peridotites (Seyler *et al.*, 2001, 2007; Morishita *et al.* 2007). New interstitial clinopyroxene grains indicate that the penetration of melt can cause reactions with rock (Pearce *et al.* 2000; Tamura *et al.*



Figure 10. (Colour online) (a) Th versus Ce and (b) La versus Nb diagrams. Global abyssal peridotite field is from Niu (2004). (c) Th versus Gd/Lu ratio. Most of the samples show similar trends of melt/rock interaction. (d) Sm/Yb versus La/Sm diagram for the HSP rocks. The yellow field shows serpentinized peridotites from refertilized protoliths (Deschamps *et al.* 2013). (e) Nb versus Pb diagram. Most of the samples show trends of magmatic processes. The fields of partial melting, melt refertilization, water–rock interaction and magmatic processes are from Chen *et al.* (2013).

2008). Based on these textural and compositional data, we suggest that the HSP rocks are related by an increasing degree of depletion due to both partial melting and melt-rock reactions.

8.d. Trace element modelling of melting

We constrain the depleted nature of the HSP rocks based on evidence from geochemical data indicating that they represent mantle residues produced by various degrees and multiple episodes of partial melting. Mysen & Kushiro (1977) argued that if the mantle undergoes 2% partial melting, its Cr# value will be 0.2, whereas with 10% partial melting, its Cr# value will reach more than 0.3. If the Cr# of spinel reaches 0.22, the host rock underwent 10% partial melting,

whereas if the Cr# value reaches 0.5, the host rock underwent 15% partial melting in the spinel lherzolite facies (Jaques & Green, 1980). Based on this idea, the Cr# values of the HSP rocks range from 0.41 to 0.52, which is consistent with 15-20% partial melting of fertile spinel lherzolite. In addition, the abundances of Ni and Yb are less sensitive to melt-rock interaction or mantle metasomatism (Parkinson & Pearce, 1998; Green et al. 2000; Deschamps et al. 2010; Uysal et al. 2012) and are useful tools for estimating the degree of partial melting (Uysal et al. 2012; Kapsiotis, 2014). On the Yb versus Ni/Yb diagram, the HSP rocks are closely related to the residue produced by the 12-25% partial melting of the primitive mantle (Fig. 12a). Chondrite-normalized REE patterns and curve calculated for non-modal fractional



Figure 11. (Colour online) Chondrite-normalized REE abundances in the HSP rocks (Sun & McDonough, 1989). The hydrothermal sediment data from the MAR (Rainbow and Saldanha serpentinite host vent) are from Dias & Barriga (2006); data for abyssal peridotites that interacted with melt (MRI) and fluids (FRI) are from Paulick *et al.* (2006); and seawater data are from Turekian (1968).

melting of depleted MORB mantle (DMM) are shown in Figure 12b; in this diagram the HREEs variation coincides with the pattern calculated for 10–25 % melt extraction (Fig. 12b), because HREEs are less mobile than LREEs during mantle metasomatism (Aldanmaz *et al.* 2009; Niu *et al.* 2015; Marchesi *et al.* 2016). The LREE abundances are significantly different from the melting model estimation and are higher than the calculated LREE contents, indicating that these mantle rocks were enriched in LREEs during the syn- and/or post-melting interaction with fluid/melts (Aldanmaz *et al.* 2009; Marchesi *et al.* 2009, 2016; Niu *et al.* 2015).

8.d. Geodynamic implications and synthesis

The Harsin–Sahneh ophiolite complex, which is Jurassic to Late Cretaceous in age, has been interpreted as a suture zone between the Arabian and Iranian Plates. Previous data (Delaloye & Desmons, 1980; Desmons & Beccaluva, 1983; Ghazi & Hassanipak, 1999; Allahyari *et al.* 2010; Saccani *et al.* 2013) have suggested that the Harsin–Sahneh ophiolite formed at a mid-ocean ridge or island arc basin, whereas others (Wrobel Daveau *et al.* 2010; Azizi *et al.* 2013; Saccani *et al.* 2013; Saccani *et al.* 2013; Saccani *et al.* 2013; Nouri *et al.* 2016, 2017) have proposed that this zone formed in an undeveloped oceanic basin that was affected by a mantle plume.



Figure 12. (Colour online) (a) Yb versus Ni/Yb diagram for the HSP rocks. The closed system non-modal dynamic melting model of primitive mantle with spinel-facies mantle mineralogy $Ol_{0.53(0.06)} + Opx_{0.27(0.67)} + Sp_{0.03(0.11)}$ is from Palme & O'Neill (2004) and Kinzler (1997). Partition coefficients are compiled from Suhr (1999), Green et al. (2000) and Adam & Green (2006). (b) Chondrite-normalized REE patterns for the HSP rocks. The samples are compared to the REE patterns calculated for anhydrous non-modal fractional melting of depleted mantle (DMM) (Salters & Stracke, 2004) in the spinel lherzolite facies (yellow dashed lines). Source and melting for olivine, orthopyroxene and clinopyroxene modal proportion are 0.54:0.28:0.18 and 0.11:0.59:0.53 at 1.5 GPa (Niu, 1997). Labels indicate partial melting degree. Partition coefficients are from Bedini & Bodinier (1999), Donnelly et al. (2004) and Su & Langmuir (2003). Normalizing values are from Sun & McDonough (1989).

As discussed above, our study shows that the HSP rocks have geochemical affinities to both MORB and supra-subduction-type rocks. These geochemical characteristics are directly related to a back-arc tectonic setting, which is supported by the elemental and isotopic signatures of the HSP rocks. We concluded that the evolution of the HSP rocks involved multiple episodes of depletion and refertilization. The less depleted HSP rocks may represent the restites produced by low degrees of partial melting during the first stage of partial melting in an extensional spreading basin. The trace element compositions of the less depleted HSP samples can be modelled by the low to moderate partial melting of a mantle source followed by interactions between the residual and percolating melt. In contrast, the highly depleted HSP rocks formed during the second stage of partial melting, followed by melt-peridotite interactions with the previously depleted mantle in an SSZ setting. The trace element abundances of the highly depleted peridotite can be explained by the high degree of partial melting of an already depleted mantle; this stage requires plate convergence and subduction, where the fluxing of the mantle by subduction melts caused the further melting of depleted peridotites to produce highly refractory residual compositions. These melting events are related to the drifting of the Arabian Plate during the Late Triassic – Early Jurassic (Delaloye & Desmons, 1980; Desmons & Beccaluva, 1983; Ghazi & Hassanipak, 1999; Dilek & Ahmed, 2003; Alavi, 2004; Jassim & Goff, 2006; Allahyari et al. 2010; Azizi et al. 2013; Saccani et al. 2013; Shafaii Moghadam & Stern, 2015; Nouri et al. 2016, 2017), such as the Neo-Tethys Ocean. During this stage, large volumes of HSP rocks were produced in an extensional spreading basin, as confirmed by the presence of some HSP rocks with a less depleted mantle origin. During this process, the HSP rocks were locally altered by seawater along deep faults and fractures. Then the initiation of subduction occurred from the Early Jurassic to the Late Cretaceous (Braud, 1978; Shahidi & Nazari, 1997; Azizi et al. 2013; Ao et al. 2016; Davoudian et al. 2016; Nouri et al. 2016), and they experienced a second melting episode, in which the less enriched component was derived from continental crustal materials associated with the subducted slab.

9. Conclusions

The high LREE contents and significantly positive correlations between LREEs and HFSEs in the HSP rocks reflect the refertilization of residual peridotites by porous melt flow. Their Sr isotopic ratios show that waterrock interactions may not have played an important role in enriching their LREE contents. In contrast, melt-rock interaction can be considered to be a main cause of the LREEs enrichment of these rocks. Based on our findings in this research, the HSP rocks were predominantly formed by the extraction of MORBlike magma from mantle peridotite at a mid-oceanic ridge or in a back-arc basin due to a combination of melt/rock interactions and a high degree of partial melting in a supra-subduction zone from the Late Triassic to Late Cretaceous.

Acknowledgements. This research is supported by a JSPS KAKENHI grant in Japan (no. 25303029). We would like to thank Nagoya University (Japan) for supporting F. Nouri in summer 2016 and the University of Kurdistan for supporting field trips. The authors would like to give special thanks to Prof. K. Yamamoto from Nagoya University for supporting our XRF and ICP-MS analyses. We are grateful to Prof. S. Wallis from Nagoya University for supporting EMP analyses and to K. Mano from Nagoya University for his technical assistance. This paper benefited greatly from critical comments by Dr C. Deering, Editor, and two anonymous reviewers.

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